Particle Dark Matter in the galactic halo: results from DAMA/LIBRA

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The Dark Side of the Universe: experimental evidences ...

COMA Cluster

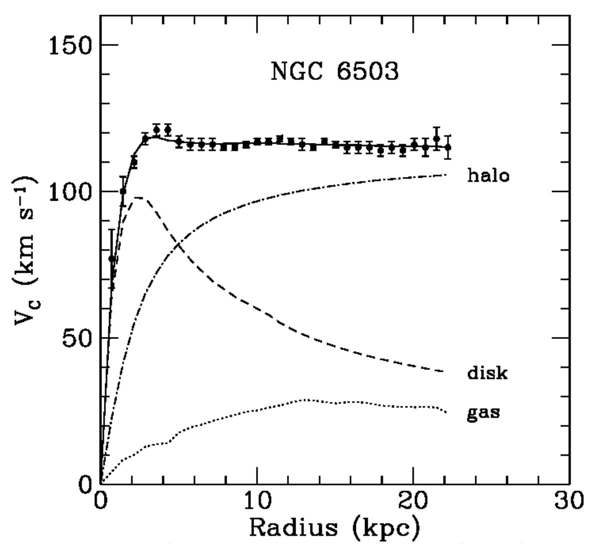
First evidence and confirmations:

1933 **F. Zwicky:** studying dispersion velocity of Coma galaxies

1936 **S. Smith:** studying the Virgo cluster

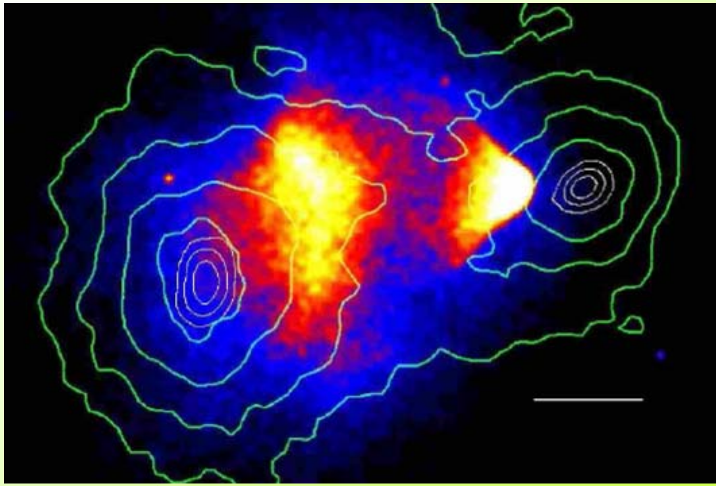
1974 **two groups:** systematical analysis of *mass density* vs *distance from center* in many galaxies

Other experimental evidences



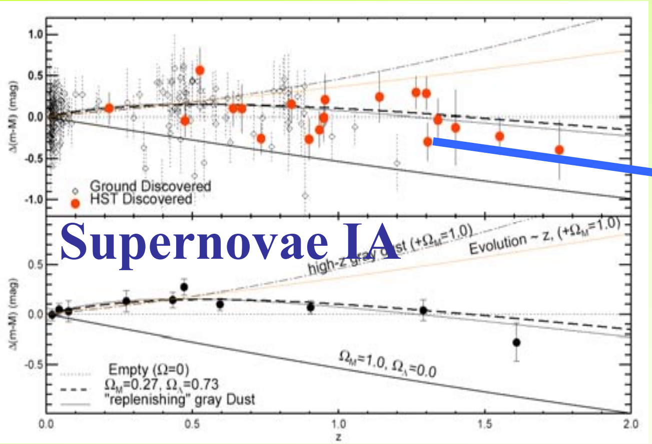
Rotational curve of a spiral galaxy

- ✓ from LMC motion around Galaxy
- ✓ from X-ray emitting gases surrounding elliptical galaxies
- ✓ from hot intergalactic plasma velocity distribution in clusters
- ✓ ...
- ✓ bullet cluster 1E0657-558

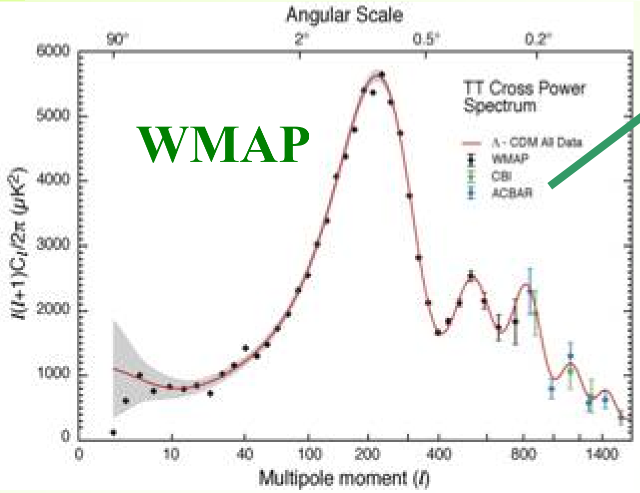


$M_{\text{visible Universe}} << M_{\text{gravitational effect}}$ ⇒ about 90% of the mass is DARK

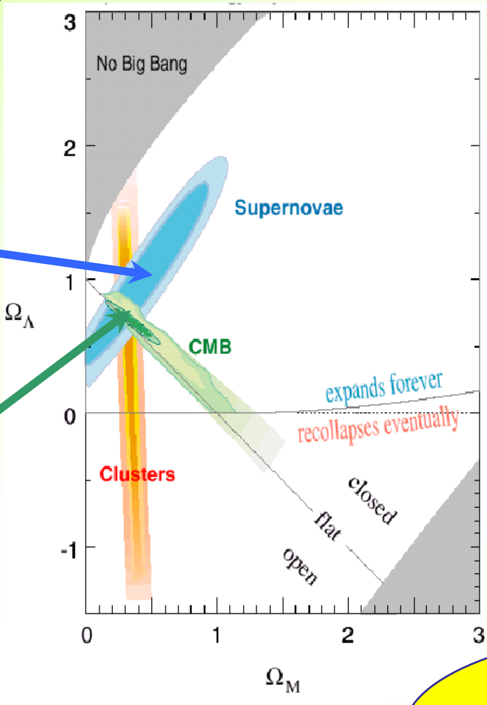
"Concordance model"



Supernovae IA



WMAP



$$\Omega = \Omega_\Lambda + \Omega_M =$$

close to 1

Ω = density/critical density

6 atoms of H/m³

$$\Omega_\Lambda \approx 0.74$$
$$\Omega_M \approx 0.26$$

The Universe is flat

Primordial Nucleosynthesis

Observations on:
- light nuclei abundance
- microlensings
- visible light.

Structure formation in the Universe

The baryons give "too small" contribution

$\Omega_b \sim 4\%$

Non baryonic **Cold Dark Matter** is dominant

$\Omega_{CDM} \sim 22\%,$
$\Omega_{HDM,\nu} < 1\ \%$

~ 90% of the matter in the Universe is non baryonic
A large part of the Universe is in form of non baryonic Cold Dark Matter particles

Relic DM particles from primordial Universe

SUSY
(as neutralino or sneutrino
In various scenarios)

the sneutrino in the Smith
and Weiner scenario

sterile ν

electron interacting dark matter

a heavy ν of the 4-th family

even a suitable particle not
yet foreseen by theories

axion-like (light pseudoscalar
and scalar candidate)

self-interacting dark matter

mirror dark matter

Kaluza-Klein particles (LKK)

heavy exotic canditates, as
"4th family atoms", ...

Elementary Black holes,
Planckian objects,
Daemons

etc…

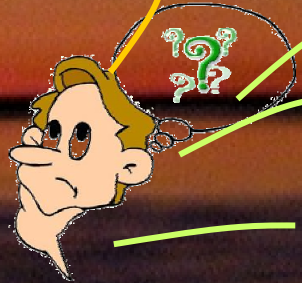
(& invisible axions, ν's)

&

Right halo model and parameters?

- **Composition?**
DM multicomponent also
in the particle part?

- **Right related nuclear and**
particle physics?

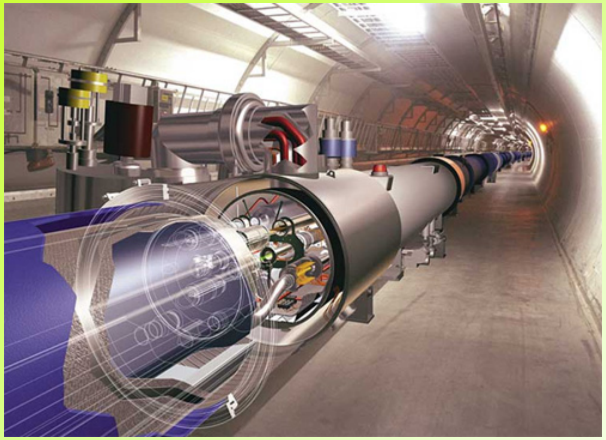


Non thermalized components?

Caustics?

clumpiness?

etc… etc…



What accelerators can do:

to demostrate the existence of some of the possible DM candidates

What accelerators cannot do:

to credit that a certain particle is the Dark Matter solution or the “single” Dark Matter particle solution…

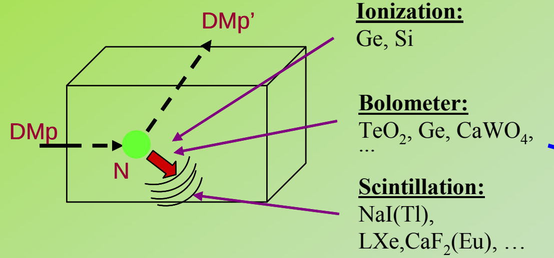
+ DM candidates and scenarios exist (even for neutralino candidate) on which accelerators cannot give any information

DM direct detection method using a model independent approach and a low-background widely-sensitive target material

Some direct detection processes:

- Scattering**s** on nuclei

  → detection of nuclear recoil energy



Ionization:
Ge, Si

Bolometer:
TeO$_2$, Ge, CaWO$_4$, …

Scintillation:
NaI(Tl),
LXe,CaF$_2$(Eu), …

- Inelastic Dark Matter: **W + N → W* + N**

  → W has Two mass states χ+ , χ- with δ mass splitting

  → Kinematical constraint for the inelastic scattering of χ- on a nucleus

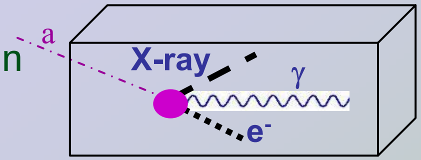
$$\frac{1}{2}\mu v^2 \ge \delta \Leftrightarrow v \ge v_{thr} = \sqrt{\frac{2\delta}{\mu}}$$

- Excitation of bound electrons in scatterings on nuclei

  → detection of recoil nuclei + e.m. radiation

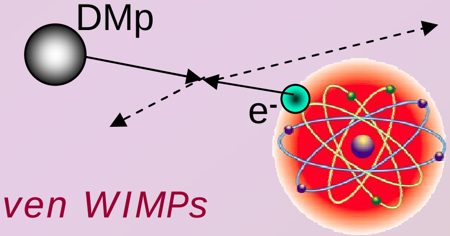
- Conversion of particle into e.m. radiation

  → detection of γ, X-rays, e$^-$



- Interaction only on atomic electrons

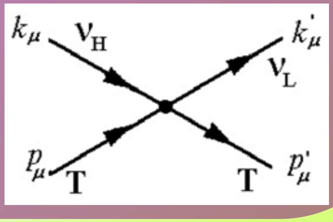
  → detection of e.m. radiation



… even WIMPs

- Interaction of ligth DMp (LDM) on e$^-$ or nucleus with production of a lighter particle

  → detection of electron/nucleus recoil energy



e.g. sterile ν

e.g. signals from these candidates are completely lost in experiments based on "rejection procedures" of the e.m. component of their rate

… also other possibilities…

- **… and more**

The direct detection experiments can be classified in two classes, depending on what they are based:

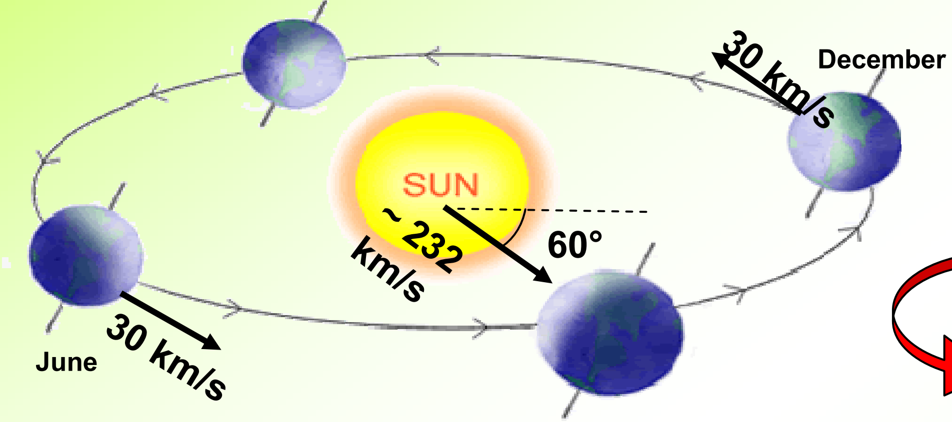
1. on the recognition of the signals due to Dark Matter particles with respect to the background by using a "model-independent" signature

2. on the use of uncertain techniques of rejection of electromagnetic background (adding systematical effects and lost of candidates with part or pure electromagnetic productions)

The annual modulation: a model independent signature for the investigation of Dark Matter particles component in the galactic halo

With the present technology, the annual modulation is the main model independent signature for the DM signal. Although the modulation effect is expected to be relatively small **a suitable large-mass, low-radioactive set-up with an efficient control of the running conditions would point out its presence.**

Drukier, Freese, Spergel PRD86
Freese et al. PRD88



- v_{sun} ~ 232 km/s (Sun velocity in the halo)
- v_{orb} = 30 km/s (Earth velocity around the Sun)
- $\gamma = \pi/3$
- $\omega = 2\pi/T$ T = 1 year
- t_0 = 2nd June (when $v_\oplus$ is maximum)

$$v_\oplus(t) = v_{sun} + v_{orb}\cos\gamma\cos[\omega(t-t_0)]$$

$$S_k[\eta(t)] = \int_{\Delta E_k} \frac{dR}{dE_R}\, dE_R \cong S_{0,k} + S_{m,k}\cos[\omega(t-t_0)]$$

Expected rate in given energy bin changes because the annual motion of the Earth around the Sun moving in the Galaxy

Requirements of the annual modulation

1) Modulated rate according cosine
2) In a definite low energy range
3) With a proper period (1 year)
4) With proper phase (about 2 June)
5) For single hit events in a multi-detector set-up
6) With modulation amplitude in the region of maximal sensitivity must be <7% for usually adopted halo distributions, but it can be larger in case of some possible scenarios

To mimic this signature, spurious effects and side reactions must not only - obviously - be able to account for the whole observed modulation amplitude, but also to satisfy contemporaneously all the requirements

Competitiveness of NaI(Tl) set-up

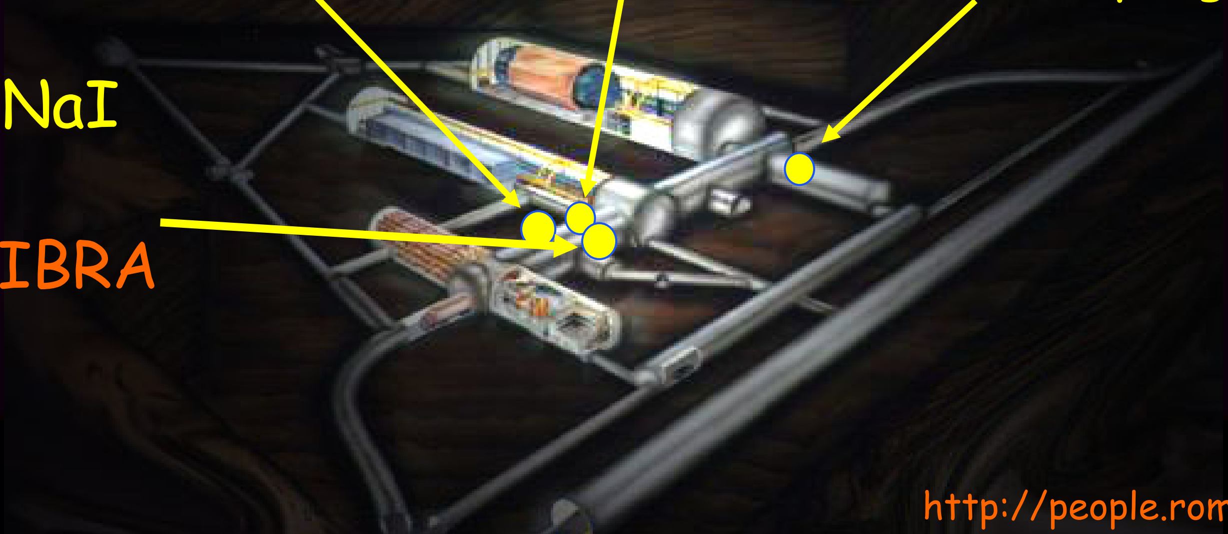
- **Well known technology**
- **High duty cycle**
- **Large mass possible**
- ***"Ecological clean"*** **set-up; no safety problems**
- **Cheaper than every other considered technique**
- **Small underground space needed**
- **High radiopurity by selections, chem./phys. purifications, protocols reachable**
- **Well controlled operational condition feasible**
- **Routine calibrations feasible down to keV range in the same conditions as the production runs**
- **Neither re-purification procedures nor cooling down/warming up (reproducibility, stability, ...)**
- **Absence of microphonic noise + effective noise rejection at threshold (τ of NaI(Tl) pulses hundreds ns, while τ of noise pulses tens ns)**
- **High light response (5.5 -7.5 ph.e./keV)**
- **Sensitive to SI, SD, SI&SD couplings and to other existing scenarios, on the contrary of many other proposed target-nuclei**
- **Sensitive to both high (by Iodine target) and low mass (by Na target) candidates**
- **Effective investigation of the annual modulation signature feasible in all the needed aspects**
- **Fragmented set-up**
- **etc.**

A low background NaI(Tl) also allows **the study of several other rare processes such as: possible processes violating the Pauli exclusion principle, CNC processes in ^{23}Na and ^{127}I, electron stability, nucleon and di-nucleon decay into invisible channels, neutral SIMP and nuclearites search, solar axion search, ...**

High benefits/cost

Roma2,Roma1,LNGS,IHEP/Beijing

DAMA: an observatory for rare processes @LNGS



DAMA/R&D

DAMA/LXe

low bckg DAMA/Ge
for sampling meas.

DAMA/NaI

DAMA/LIBRA

http://people.roma2.infn.it/dama

DAMA/LXe: results on rare processes

Dark Matter Investigation

NIMA482(2002)728

Limits on recoils investigating the DMp-^{129}Xe elastic scattering by means of PSD	PLB436(1998)379
Limits on DMp-^{129}Xe inelastic scattering	PLB387(1996)222, NJP2(2000)15.1
Neutron calibration	PLB436(1998)379, EPJdirectC11(2001)1
^{129}Xe *vs* ^{136}Xe by using PSD → SD *vs* SI signals to increase the sensitivity on the SD component	foreseen/in progress

Other rare processes:

Electron decay into invisible channels	Astrop.P.5(1996)217
Nuclear level excitation of ^{129}Xe during CNC processes	PLB465(1999)315
N, NN decay into invisible channels in ^{129}Xe	PLB493(2000)12
Electron decay: $e^- \to \nu_e \gamma$	PRD61(2000)117301
2β decay in ^{136}Xe	Xenon01
2β decay in ^{134}Xe	PLB527(2002)182
Improved results on 2β in ^{134}Xe,^{136}Xe	PLB546(2002)23
CNC decay ^{136}Xe → ^{136}Cs	Beyond the Desert (2003) 365
N, NN, NNN decay into invisible channels in ^{136}Xe	EPJA27 s01 (2006) 35

DAMA/R&D set-up: results on rare processes

- Particle Dark Matter search with CaF_2(Eu) — NPB563(1999)97, Astrop.Phys.7(1997)73

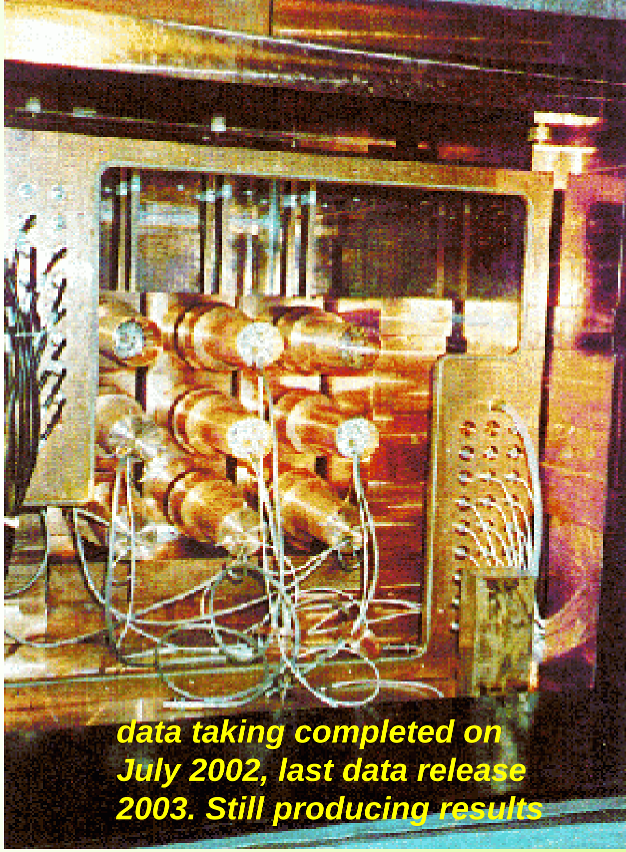
2β decay in ^{136}Ce and in ^{142}Ce	Il Nuov.Cim.A110(1997)189
2EC2ν ^{40}Ca decay	Astrop. Phys. 7(1997)73
2β decay in ^{46}Ca and in ^{40}Ca	NPB563(1999)97
$2\beta^+$ decay in ^{106}Cd	Astrop.Phys.10(1999)115
2β and β decay in ^{48}Ca	NPA705(2002)29
2EC2ν in ^{136}Ce, in ^{138}Ce and α decay in ^{142}Ce	NIMA498(2003)352
$2\beta^+$ 0ν and EC β^+ 0ν decay in ^{130}Ba	NIMA525(2004)535
Cluster decay in $LaCl_3$(Ce)	NIMA555(2005)270
CNC decay ^{139}La → ^{139}Ce	UJP51(2006)1037
α decay of natural Eu	NPA789(2007)15
β decay of ^{113}Cd	PRC76(2007)064603
$\beta\beta$ decay of ^{64}Zn	PLB658(2008)193
$\beta\beta$ decay of ^{108}Cd and ^{114}Cd	EPJA36(2008)167

DAMA/Ge & LNGS Ge facility

- RDs on highly radiopure NaI(Tl) set-up;
- several RDs on low background PMTs;
- qualification of many materials
- measurements with a $Li_6Eu(BO_3)_3$ crystal (NIMA572(2007)734)
- measurements with ^{100}Mo sample investigating $\beta\beta$ decay in the 4π low-bckg HP Ge facility of LNGS (to appear on Nucl. Phys. and Atomic Energy)
- search for ^{7}Li solar axions (NPA806(2008)388)

+Many other meas. already scheduled for near future

DAMA/NaI : ≈100 kg NaI(Tl)



data taking completed on July 2002, last data release 2003. Still producing results

Performances: N.Cim.A112(1999)545-575, EPJC18(2000)283, Riv.N.Cim.26 n. 1(2003)1-73, IJMPD13(2004)2127

Results on rare processes:

Possible Pauli exclusion principle violation	PLB408(1997)439
CNC processes	PRC60(1999)065501
Electron stability and non-paulian transitions in Iodine atoms (by L-shell)	PLB460(1999)235
Search for solar axions	PLB515(2001)6
Exotic Matter search	EPJdirect C14(2002)1
Search for superdense nuclear matter	EPJA23(2005)7
Search for heavy clusters decays	EPJA24(2005)51

Results on DM particles:

PSD	PLB389(1996)757
Investigation on diurnal effect	N.Cim.A112(1999)1541
Exotic Dark Matter search	PRL83(1999)4918

- Annual Modulation Signature

PLB424(1998)195, PLB450(1999)448, PRD61(1999)023512, PLB480(2000)23, EPJC18(2000)283, PLB509(2001)197, EPJC23(2002)61, PRD66(2002)043503, Riv.N.Cim.26 n.1 (2003)1, IJMPD13(2004)2127, IJMPA21(2006)1445, EPJC47(2006)263, IJMPA22(2007)3155, EPJC53(2008)205, PRD77(2008)023506, MPLA23(2008)2125.

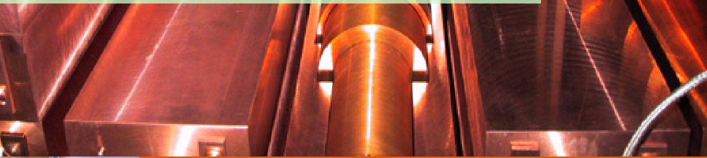
model independent evidence of a particle DM component in the galactic halo at 6.3σ C.L.

total exposure (7 annual cycles) 0.29 ton x yr

The new DAMA/LIBRA set-up ~250 kg NaI(Tl)
(Large sodium Iodide Bulk for RAre processes)

As a result of a second generation R&D for more radiopure NaI(Tl)
by exploiting new chemical/physical radiopurification techniques
(all operations involving crystals and PMTs - including photos - in HP Nitrogen atmosphere)

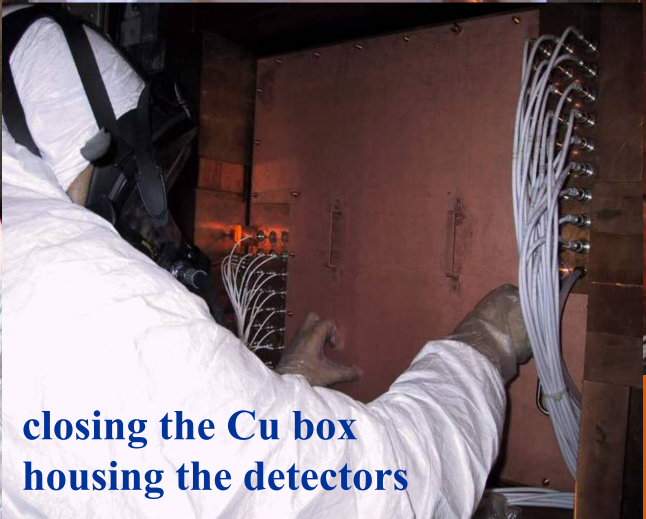
installing DAMA/LIBRA detectors



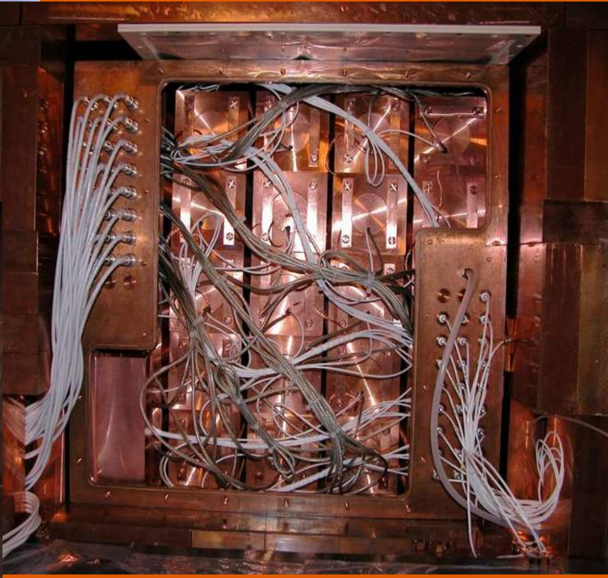
detectors during installation; in the central and right up detectors the new shaped Cu shield surrounding light guides (acting also as optical windows) and PMTs was not yet applied

assembling a DAMA/ LIBRA detector

filling the inner Cu box with further shield



closing the Cu box housing the detectors



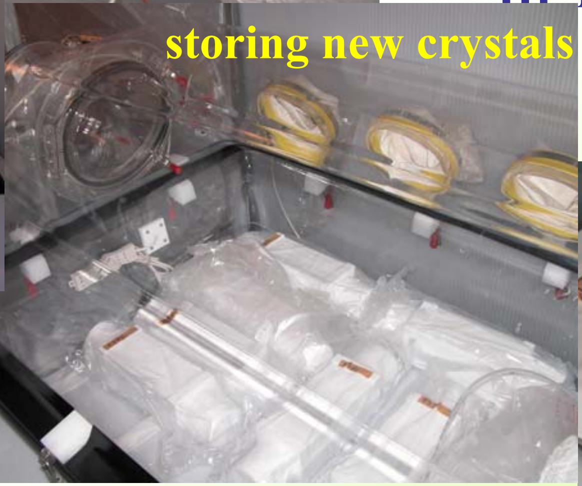
view at end of detectors' installation in the Cu box

DAMA/LIBRA ~250 kg NaI(Tl)
(Large sodium Iodide Bulk for RAre processes)

As a result of a second generation R&D for more radiopure NaI(Tl)
by exploiting new chemical/physical radiopurification techniques
(all operations involving crystals and PMTs - including photos - in HP Nitrogen atmosphere)

improving installation
and environment

PMT
+HV
divider



storing new crystals

Cu etching with super- and ultra-pure HCl solutions, dried and sealed in HP N_2

etching staff at work
in clean room

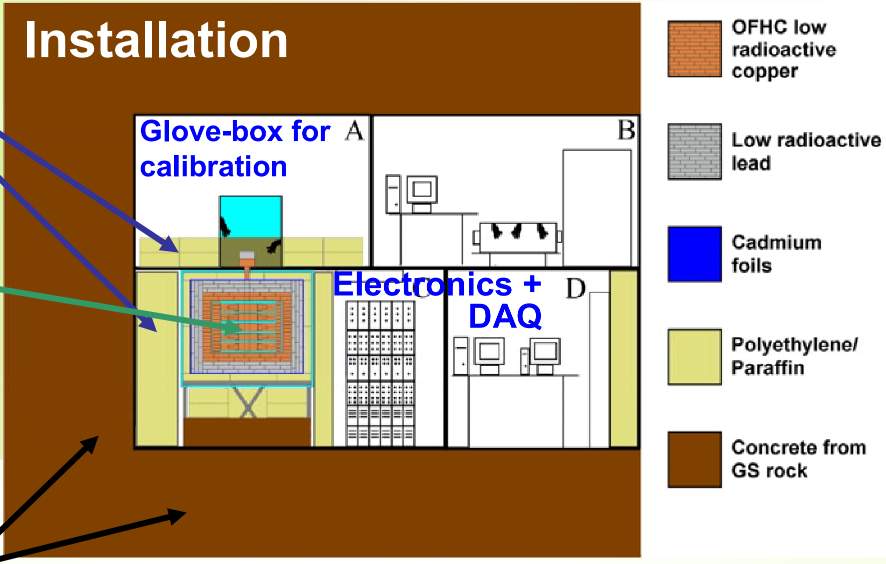
The DAMA/LIBRA set-up

For details, radiopurity,performances, procedures, etc.
NIMA592(2008)297

Polyethylene/
paraffin

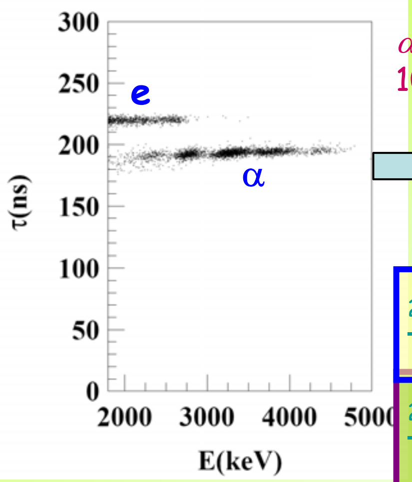
- 25 x 9.7 kg NaI(Tl) in a 5x5 matrix
- two Suprasil-B light guides directly coupled to each bare crystal
- two PMTs working in coincidence at the single ph. el. threshold

~ 1m concrete from GS rock



- Dismounting/Installing protocol (with "Scuba" system)
- All the materials selected for low radioactivity
- Multicomponent passive shield
- Three-level system to exclude Radon from the detectors
- Calibrations in the same running conditions as production runs
- Installation in air conditioning + huge heat capacity of shield
- Monitoring/alarm system; many parameters acquired with the production data
- Pulse shape recorded by Waweform Analyzer *TVS641A* (2chs per detector), 1 *Gsample/s,* 8 bit, bandwidth 250 MHz
- Data collected from low energy up to MeV region, despite the hardware optimization was done for the low energy

Some on residual contaminants in new NaI(Tl) detectors



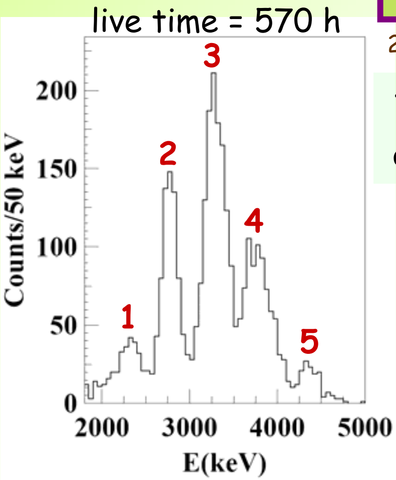
α/e pulse shape discrimination has practically 100% effectiveness in the MeV range

The measured α yield in the new DAMA/LIBRA detectors ranges from 7 to some tens α/kg/day

Second generation R&D for new DAMA/LIBRA crystals: new selected powders, physical/chemical radiopurification, new selection of overall materials, new protocol for growing and handling

^{232}Th residual contamination From time-amplitude method. If ^{232}Th chain at equilibrium: it ranges from 0.5 ppt to 7.5 ppt

^{238}U residual contamination First estimate: considering the measured α and ^{232}Th activity, if ^{238}U chain at equilibrium ⇒ ^{238}U contents in new detectors typically range from 0.7 to 10 ppt

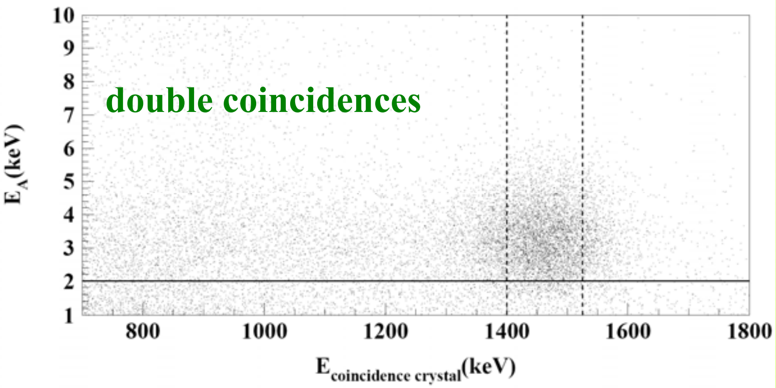


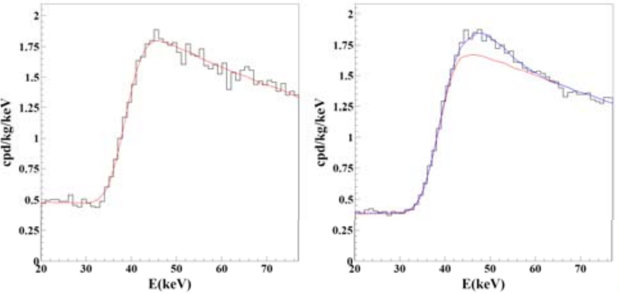
^{238}U chain splitted into 5 subchains: $^{238}U \rightarrow ^{234}U \rightarrow ^{230}Th \rightarrow ^{226}Ra \rightarrow ^{210}Pb \rightarrow ^{206}Pb$

Thus, in this case: (2.1±0.1) ppt of ^{232}Th; (0.35 ±0.06) ppt for ^{238}U
and: (15.8±1.6) μBq/kg for ^{234}U + ^{230}Th; (21.7±1.1) μBq/kg for ^{226}Ra; (24.2±1.6) μBq/kg for ^{210}Pb.

^{nat}K residual contamination

The analysis has given for the ^{nat}K content in the crystals values not exceeding about 20 ppb





^{129}I and ^{210}Pb

$^{129}I/^{nat}I \approx 1.7\times10^{-13}$ for all the new detectors

^{210}Pb in the new detectors: (5 – 30) μBq/kg.

No sizeable surface pollution by Radon daugthers, thanks to the new handling protocols

... more on NIMA592(2008)297

Some on residual contaminants in NaI(Tl) detectors

Second generation R&D for new DAMA/LIBRA crystals: new selected powders, physical/chemical radiopurification, new selection of overall materials, new protocol for growing and handling

^{232}Th residual contamination

Time-amplitude method: arrival time and energy of each event used for selection of fast decay chains in ^{232}Th family

^{224}Ra (Q_α=5.8 MeV, $T_{1/2}$=3.66 d)
→ ^{220}Rn (Q_α=6.4 MeV, $T_{1/2}$=55.6 s)
→ ^{216}Po (Q_α=6.9 MeV, $T_{1/2}$=0.145 s)
→ ^{212}Pb

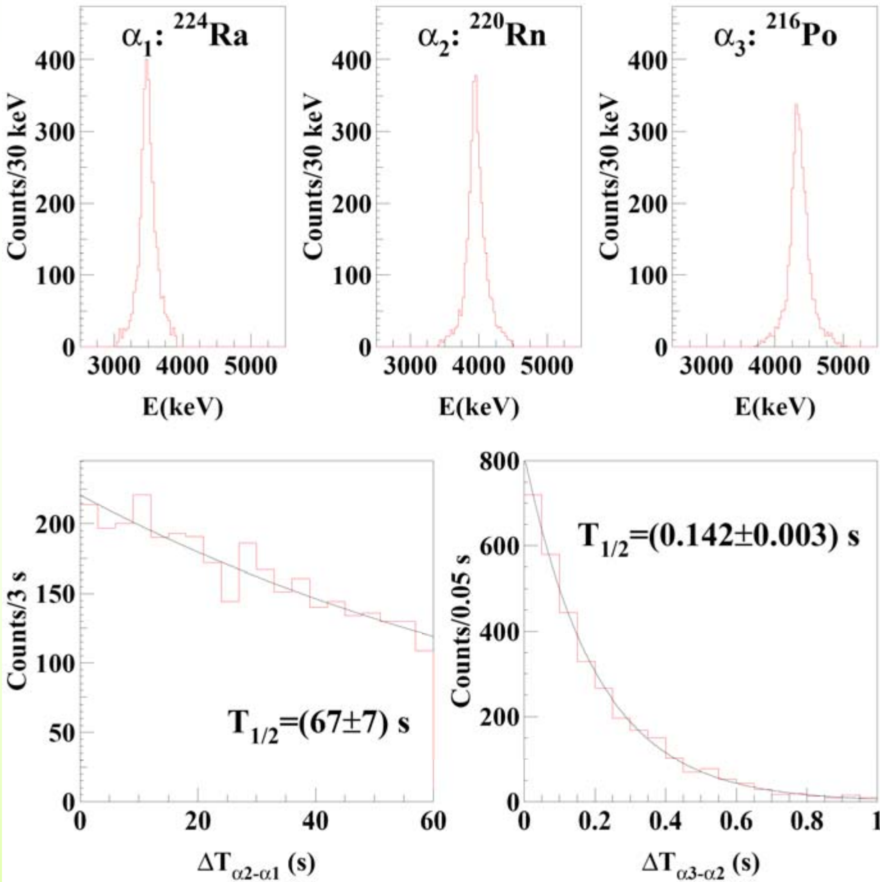
α peaks as well as the distributions of the time intervals between the events are in a good agreement with those expected

$$\alpha / \beta = 0.467(6) + 0.0257(10) \times E_\alpha [MeV]$$

⇒ ^{228}Th activity ranging from 2 to about 30 μBq/kg in the DAMA/LIBRA detectors (in agreement with Bi-Po analysis)

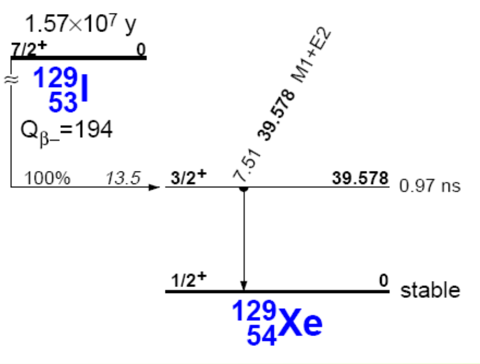
If ^{232}Th chain at equilibrium: ^{232}Th contents in new detectors typically range from 0.5 ppt to 7.5 ppt

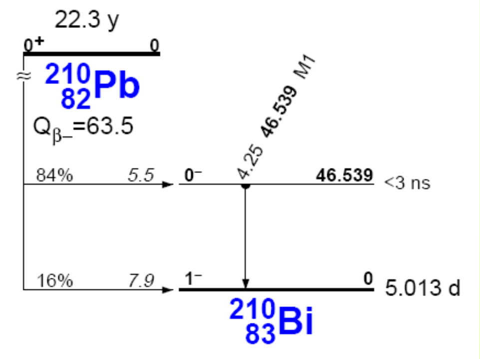
Example: 3310 triple delayed coincidences in 8100 kg×day → (9.0±0.4) μBq/kg



^{129}I and ^{210}Pb

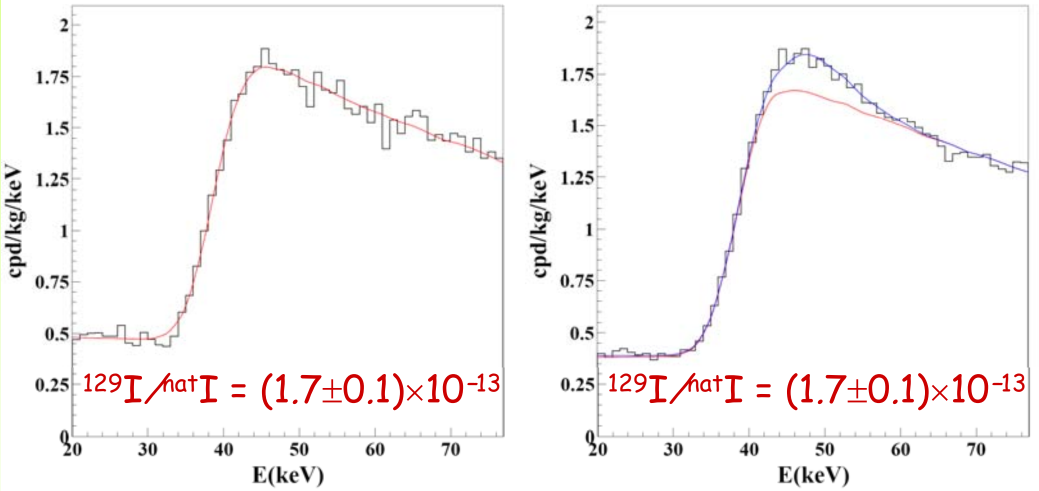
- ^{129}I ($T_{1/2}$ = 1.57×10^{7} yr) can be present in the natural Iodine with a percentage of the order of 1.5×10^{-12}
- ^{210}Pb long-lived isotope from ^{238}U chain





^{210}Pb: a= (7±4) μBq/kg

^{210}Pb: a= (29±3) μBq/kg



^{129}I/^{nat}I = (1.7±0.1)×10^{-13}

^{129}I/^{nat}I = (1.7±0.1)×10^{-13}

Energy distributions of two new DAMA/LIBRA detectors (histogram)
+ model of background for cosmogenic ^{129}I contribution (red lines)
+ contribution of internal ^{210}Pb (visible only in the second detector)

The amount of cosmogenic ^{129}I is at the same level (≈1.7×10^{-13}) for all the new detectors
(if used for dating the NaI powders ⇒ extracted from ore with an age of order of 50 Myr)

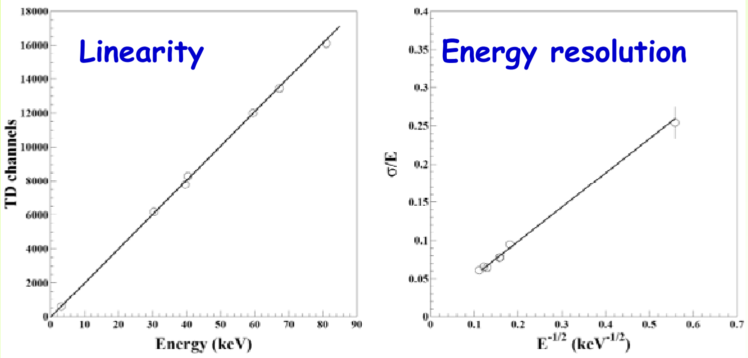
^{210}Pb in the new DAMA/LIBRA detectors typically ranges: (5 – 30) μBq/kg.

No sizeable surface pollution by Radon daugthers, thanks to the new handling protocols

... more on NIMA592(2008) 297

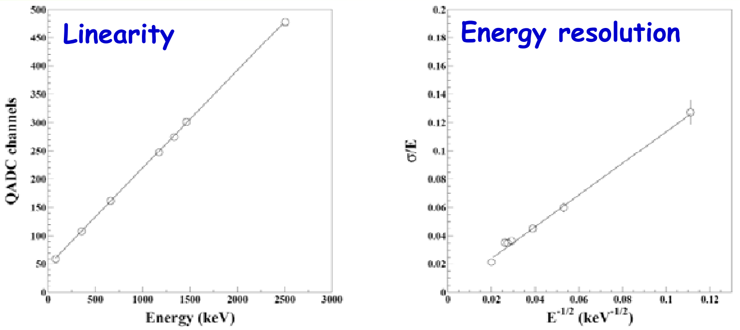
DAMA/LIBRA calibrations

Low energy: various external gamma sources (^{241}Am, ^{133}Ba) and internal X-rays or gamma's (^{40}K, ^{125}I, ^{129}I), routine calibrations with ^{241}Am

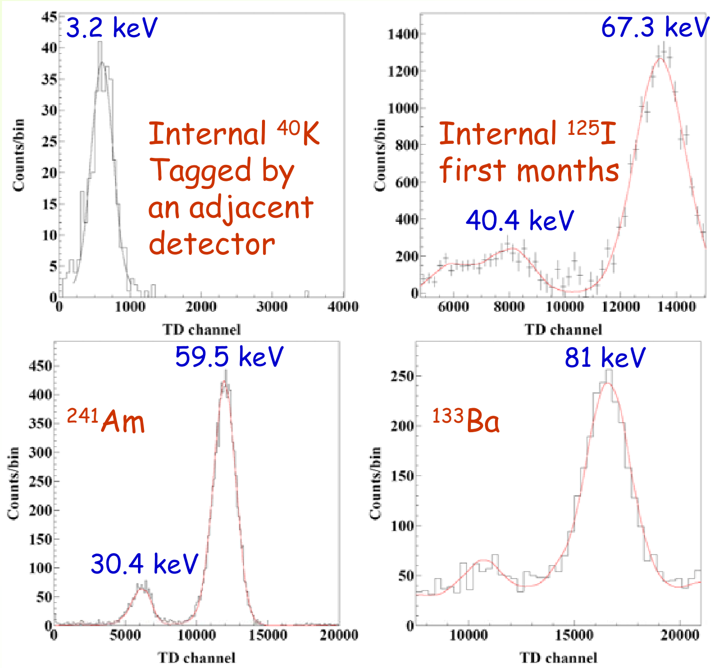


$$\frac{\sigma_{LE}}{E} = \frac{(0.448 \pm 0.035)}{\sqrt{E(keV)}} + (9.1 \pm 5.1) \cdot 10^{-3}$$

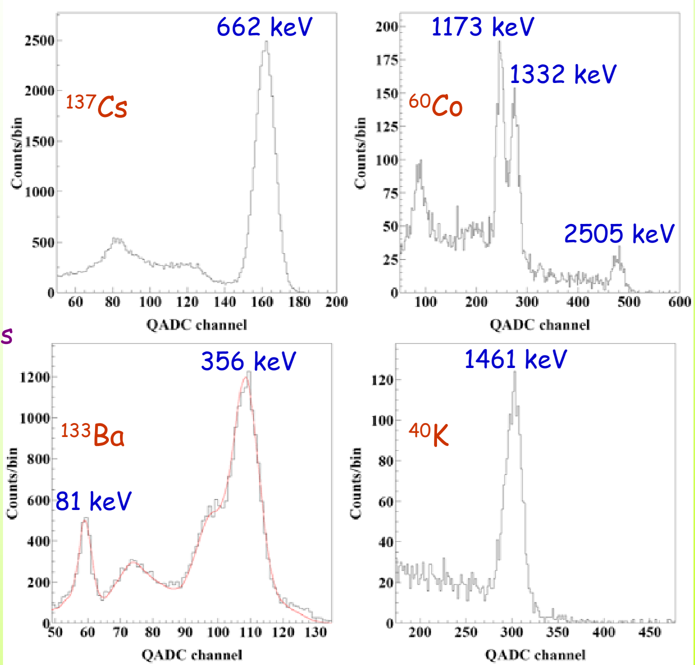
High energy: external sources of gamma rays (e.g. ^{137}Cs, ^{60}Co and ^{133}Ba) and gamma rays of 1461 keV due to ^{40}K decays in an adjacent detector, tagged by the 3.2 keV X-rays



$$\frac{\sigma_{HE}}{E} = \frac{(1.12 \pm 0.06)}{\sqrt{E(keV)}} + (17 \pm 23) \cdot 10^{-4}$$



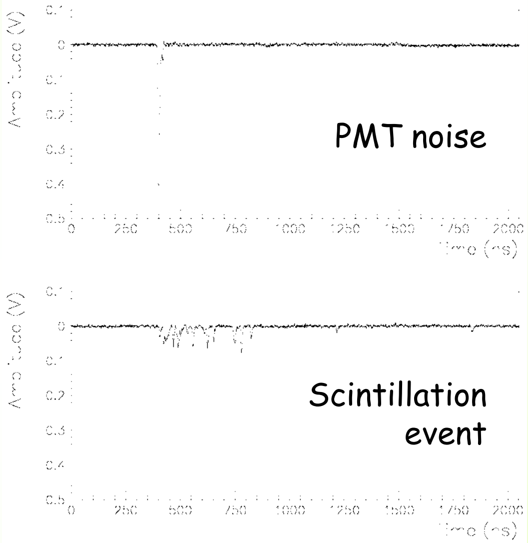
The curves superimposed to the experimental data have been obtained by simulations



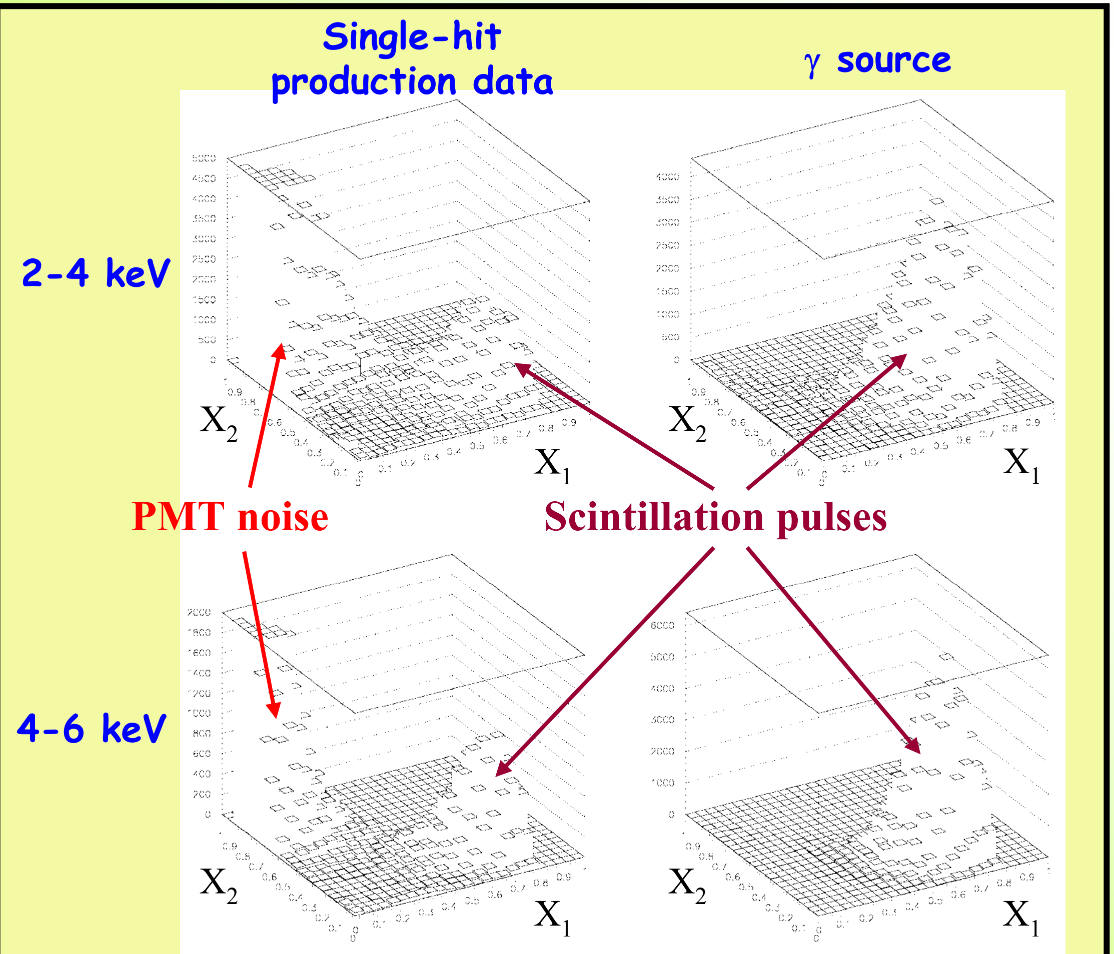
The signals (unlike low energy events) for high energy events are taken only from one PMT

Noise rejection near the energy threshold

Typical pulse profiles of PMT noise and of scintillation event with the same area, just above the energy threshold of 2 keV



The different time characteristics of PMT noise (decay time of order of tens of ns) and of scintillation event (decay time about 240 ns) can be investigated building several variables



Single-hit production data

γ source

2-4 keV

4-6 keV

PMT noise

Scintillation pulses

From the Waveform Analyser 2048 ns time window:

$$X_1 = \frac{\text{Area (from 100 ns to 600 ns)}}{\text{Area (from 0 ns to 600 ns)}};$$

$$X_2 = \frac{\text{Area (from 0 ns to 50 ns)}}{\text{Area (from 0 ns to 600 ns)}}$$

- The separation between noise and scintillation pulses is very good.
- Very clean samples of scintillation events selected by stringent acceptance windows.
- The related efficiencies evaluated by calibrations with ^{241}Am sources of suitable activity in the same experimental conditions and energy range as the production data (efficiency measurements performed each ~10 days; typically 10^4–10^5 events per keV collected)

This is the only procedure applied to the analysed data

Infos about DAMA/LIBRA data taking

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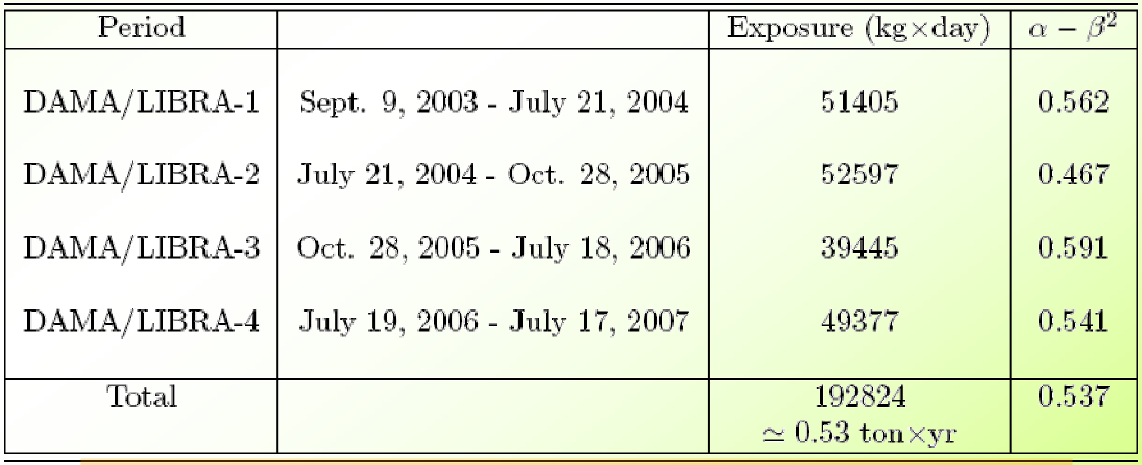
DAMA/LIBRA test runs: from March 2003 to September 2003

DAMA/LIBRA normal operation: from September 2003 to August 2004

High energy runs for TDs: September 2004
to allow internal α's identification
(approximative exposure ≈ 5000 kg × d)

DAMA/LIBRA normal operation: from October 2004



Period		Exposure (kg×day)	$\alpha - \beta^2$
DAMA/LIBRA-1	Sept. 9, 2003 - July 21, 2004	51405	0.562
DAMA/LIBRA-2	July 21, 2004 - Oct. 28, 2005	52597	0.467
DAMA/LIBRA-3	Oct. 28, 2005 - July 18, 2006	39445	0.591
DAMA/LIBRA-4	July 19, 2006 - July 17, 2007	49377	0.541
Total		192824 ≃ 0.53 ton×yr	0.537

Data released here:

- four annual cycles: 0.53 ton × yr
- calibrations: acquired ≈ 44 M events from sources
- acceptance window eff: acquired ≈ 2 M events/keV

DAMA/NaI (7 years) + DAMA/LIBRA (4 years)

total exposure: 300555 kg×day = 0.82 ton×yr

Two remarks:

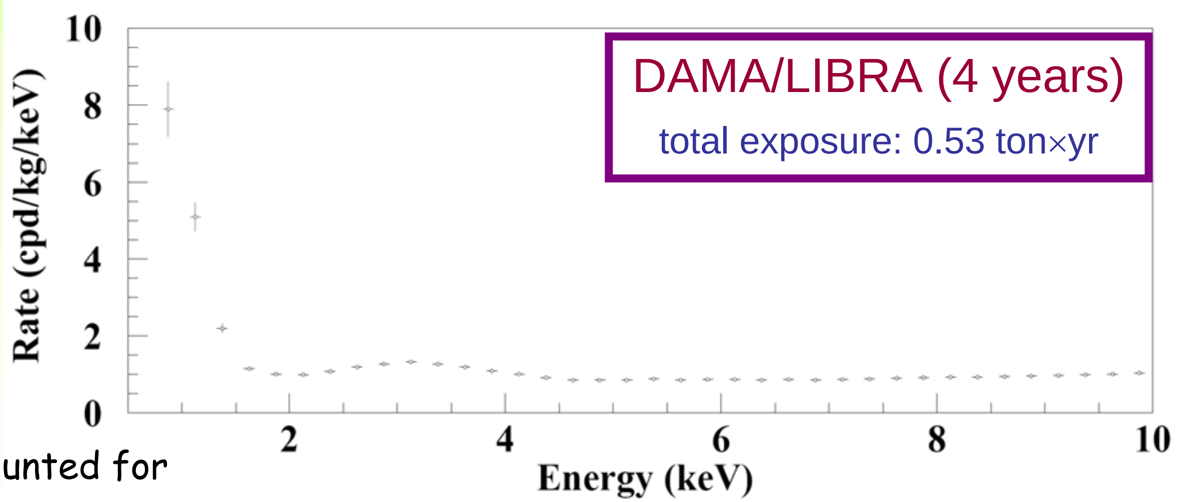
- One PMT problems after 6 months. Detector out of trigger since Sep. 2003 (since Sept. 2008 again in operation)
- Residual cosmogenic ^{125}I presence in the first year in some detectors (this motivates the Sept. 2003 as starting time)

DAMA/LIBRA is continuously running

Cumulative low-energy distribution of the *single-hit* scintillation events

Single-hit events = each detector has all the others as anticoincidence

(Obviously differences among detectors are present depending e.g. on each specific level and location of residual contaminants, on the detector's location in the 5x5 matrix, etc.)



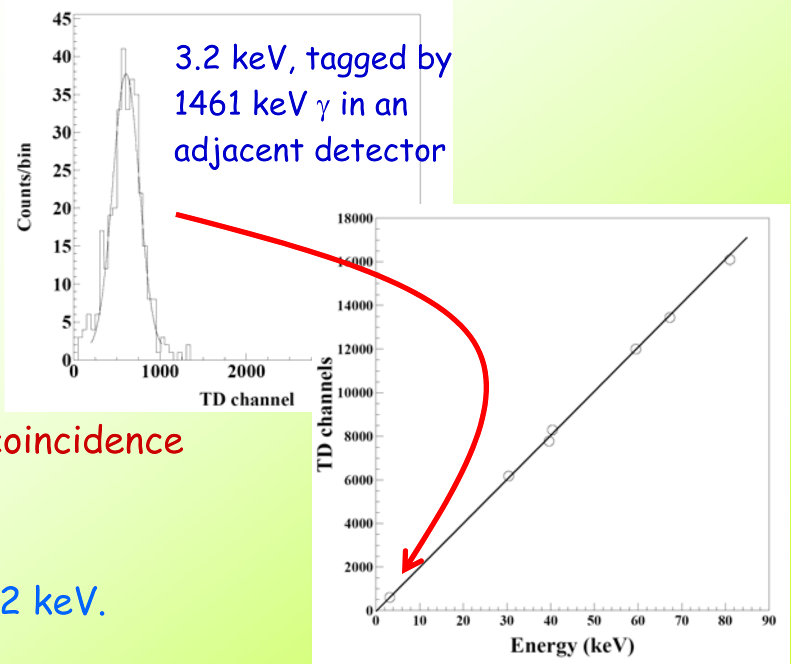
DAMA/LIBRA (4 years)

total exposure: 0.53 ton×yr

Efficiencies already accounted for

About the energy threshold:

- The DAMA/LIBRA detectors have been calibrated down to the keV region. This assures a clear knowledge of the "physical" energy threshold of the experiment.
- It obviously profits of the relatively high number of available photoelectrons/keV (from 5.5 to 7.5).
- The two PMTs of each detector in DAMA/LIBRA work in coincidence with hardware threshold at single photoelectron level.
- Effective near-threshold-noise full rejection.
- The software energy threshold used by the experiment is 2 keV.

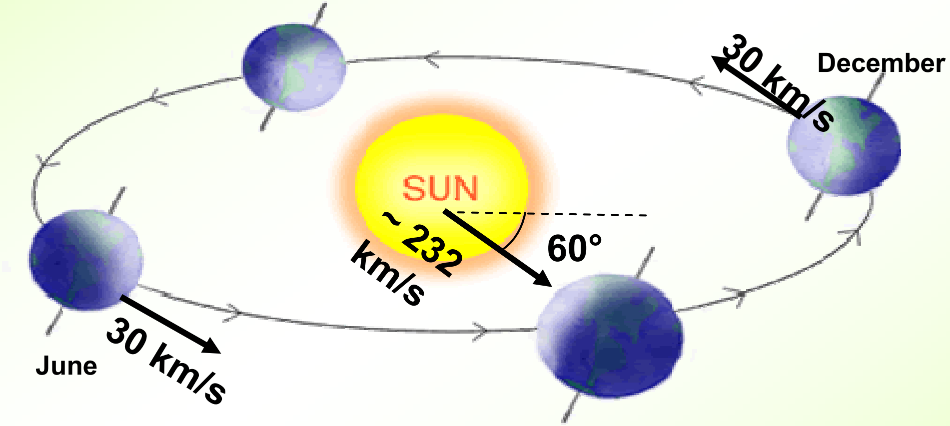


3.2 keV, tagged by 1461 keV γ in an adjacent detector

Experimental *single-hit* residuals rate vs time and energy

- Model-independent investigation of the annual modulation signature has been carried out by exploiting the time behaviour of the residual rates of the *single-hit* events in the lowest energy regions of the DAMA/LIBRA data.
- These residual rates are calculated from the measured rate of the *single-hit* events (obviously corrections for the overall efficiency and for the acquisition dead time are already applied) after subtracting the constant part:

$$\left\langle r_{ijk} - flat_{jk} \right\rangle_{jk}$$



- r_{ijk} is the rate in the considered *i-th* time interval for the *j-th* detector in the *k-th* energy bin
- $flat_{jk}$ is the rate of the *j-th* detector in the *k-th* energy bin averaged over the cycles.
- The average is made on all the detectors (*j* index) and on all the energy bins (*k* index)
- The weighted mean of the residuals must obviously be zero over one cycle.

Model Independent Annual Modulation Result

DAMA/NaI (7 years) + DAMA/LIBRA (4 years) Total exposure: 300555 kg×day = 0.82 ton×yr

experimental single-hit residuals rate vs time and energy

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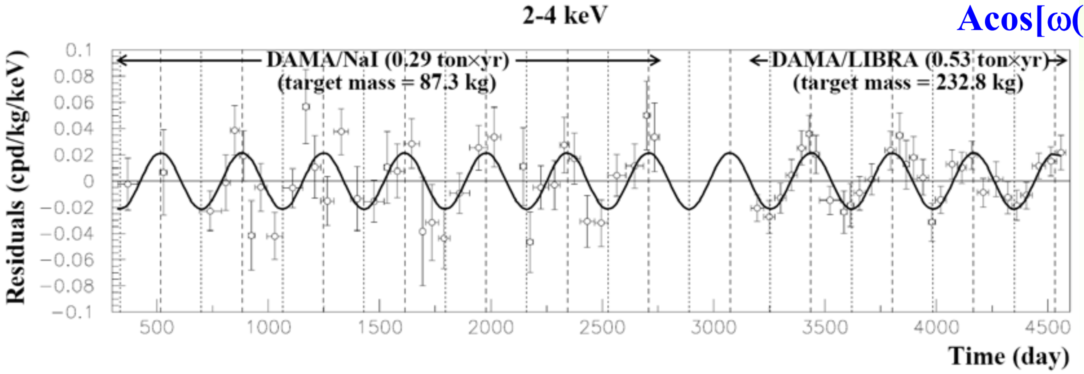
$A\cos[\omega(t-t_0)]$; continuous lines: t_0 = 152.5 d, T = 1.00 y



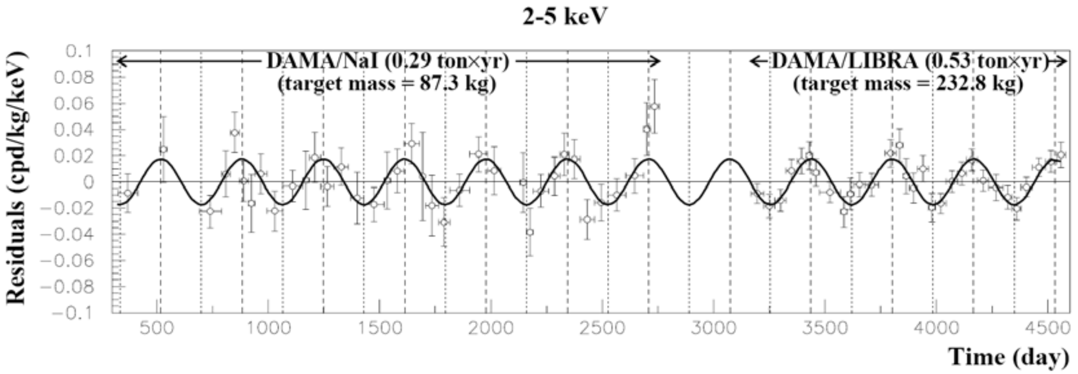
2-4 keV

A=(0.0215±0.0026) cpd/kg/keV

χ^2/dof = 51.9/66 **8.3 σ C.L.**

Absence of modulation? No

χ^2/dof=117.7/67 ⇒ P(A=0) = 1.3×10^{-4}



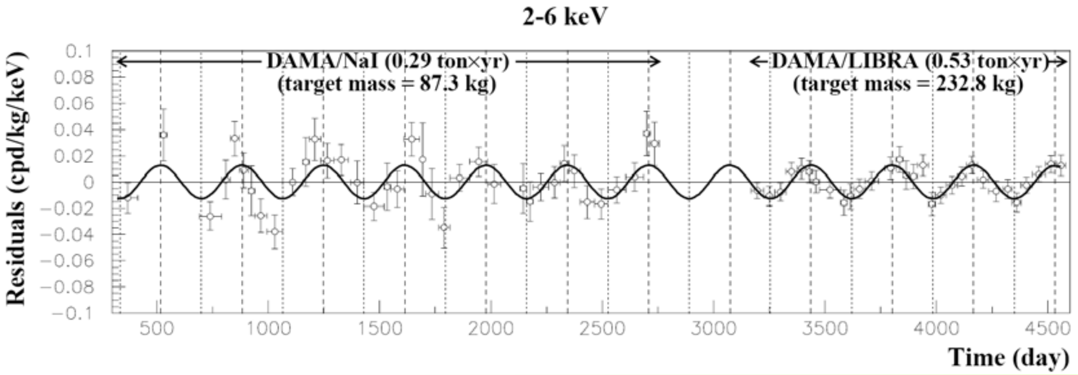
2-5 keV

A=(0.0176±0.0020) cpd/kg/keV

χ^2/dof = 39.6/66 **8.8 σ C.L.**

Absence of modulation? No

χ^2/dof=116.1/67 ⇒ P(A=0) = 1.9×10^{-4}



2-6 keV

A=(0.0129±0.0016) cpd/kg/keV

χ^2/dof = 54.3/66 **8.2 σ C.L.**

Absence of modulation? No

χ^2/dof=116.4/67 ⇒ P(A=0) = 1.8×10^{-4}

The data favor the presence of a modulated behavior with proper features at 8.2σ C.L.

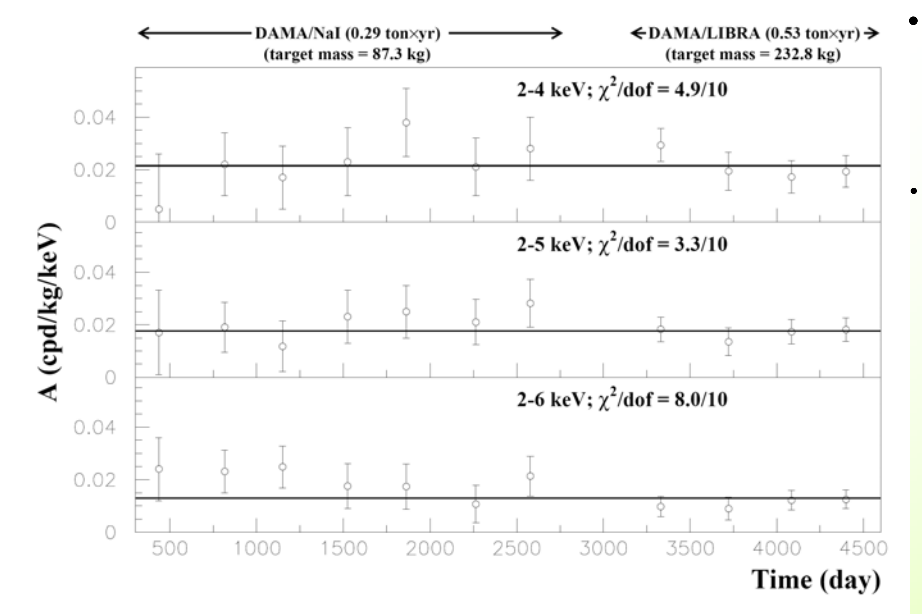
Model-independent residual rate for single-hit events

DAMA/NaI (7 years) + DAMA/LIBRA (4 years) total exposure: 300555 kg×day = 0.82 ton×yr

Results of the fits keeping the parameters free:

	A (cpd/kg/keV)	T= 2π/ω (yr)	t_0 (day)	C.L.
DAMA/NaI (7 years)				
(2÷4) keV	0.0252 ± 0.0050	1.01 ± 0.02	125 ± 30	5.0σ
(2÷5) keV	0.0215 ± 0.0039	1.01 ± 0.02	140 ± 30	5.5σ
(2÷6) keV	0.0200 ± 0.0032	1.00 ± 0.01	140 ± 22	6.3σ
DAMA/LIBRA (4 years)				
(2÷4) keV	0.0213 ± 0.0032	0.997 ± 0.002	139 ± 10	6.7σ
(2÷5) keV	0.0165 ± 0.0024	0.998 ± 0.002	143 ± 9	6.9σ
(2÷6) keV	0.0107 ± 0.0019	0.998 ± 0.003	144 ± 11	5.6σ
DAMA/NaI + DAMA/LIBRA				
(2÷4) keV	0.0223 ± 0.0027	0.996 ± 0.002	138 ± 7	8.3σ
(2÷5) keV	0.0178 ± 0.0020	0.998 ± 0.002	145 ± 7	8.9σ
(2÷6) keV	0.0131 ± 0.0016	0.998 ± 0.003	144 ± 8	8.2σ

Modulation amplitudes, *A*, of single year measured in the 11 one-year experiments of DAMA (NaI + LIBRA)



- The modulation amplitudes for the (2 – 6) keV energy interval, obtained when fixing exactly the period at 1 yr and the phase at 152.5 days, are: (0.019 ± 0.003) cpd/kg/keV for DAMA/NaI and (0.011 ± 0.002) cpd/kg/keV for DAMA/LIBRA.
- Thus, their difference: (0.008 ± 0.004) cpd/kg/keV is ≈ *2σ* which corresponds to a modest, but non negligible probability.

χ^2 test (χ^2/dof = 4.9/10, 3.3/10 and 8.0/10) and *run test* (lower tail probabilities of 74%, 61% and 11%) accept at 90% C.L. the hypothesis that the modulation amplitudes are normally fluctuating around their best fit values.

Compatibility among the annual cycles

Power spectrum of single-hit residuals

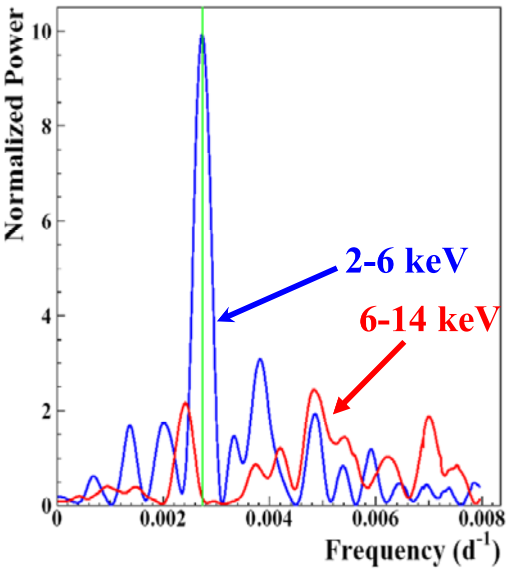
(according to Ap.J.263(1982)835; Ap.J.338(1989)277)

Treatment of the experimental errors and time binning included here

2-6 keV vs 6-14 keV

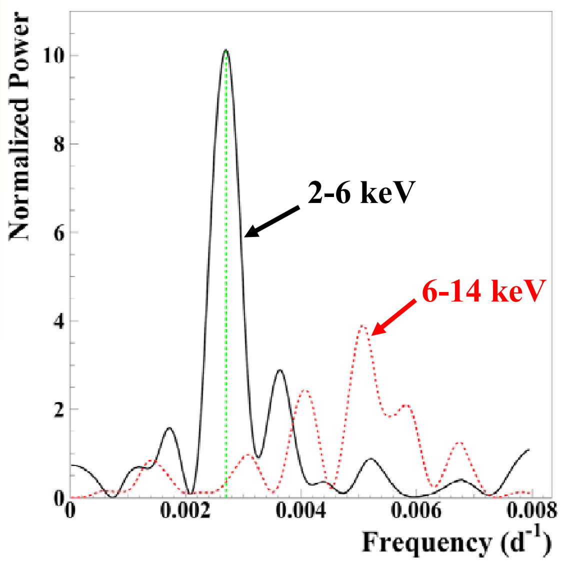
DAMA/NaI (7 years)

total exposure: 0.29 ton×yr



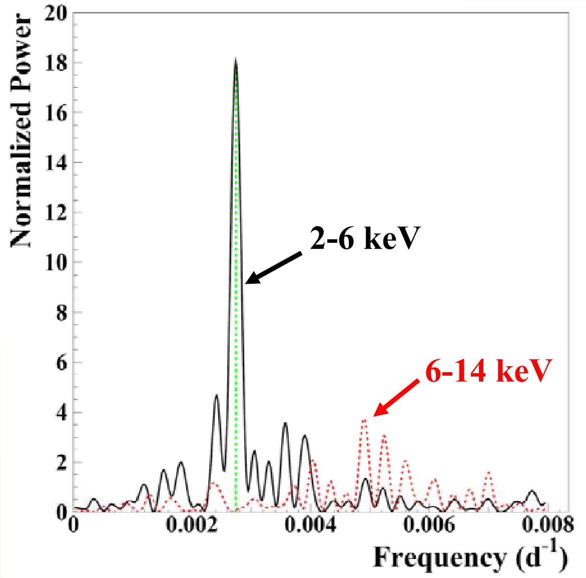
DAMA/LIBRA (4 years)

total exposure: 0.53 ton×yr



DAMA/NaI (7 years) + DAMA/LIBRA (4 years)

total exposure: 0.82 ton×yr



Principal mode in the 2-6 keV region:

DAMA/NaI	DAMA/LIBRA	DAMA/NaI+LIBRA
$2.737 \cdot 10^{-3}$ d^{-1} ≈ 1 y^{-1}	2.705×10^{-3} d^{-1} ≈ 1 yr^{-1}	2.737×10^{-3} d^{-1} ≈ 1 yr^{-1}

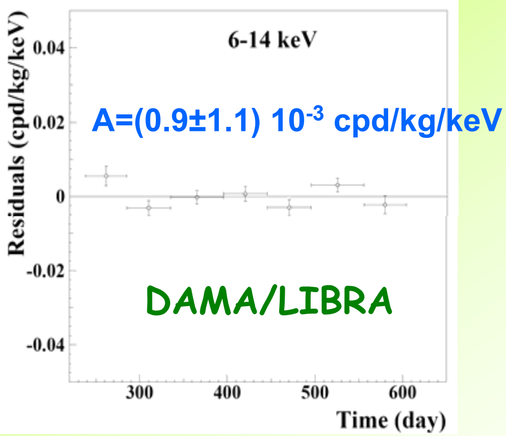
+

Not present in the 6-14 keV region (only aliasing peaks)

Clear annual modulation is evident in (2-6) keV while it is absence just above 6 keV

Can a hypothetical background modulation account for the observed effect?

- **No Modulation above 6 keV**



Mod. Ampl. (6-10 keV): cpd/kg/keV
(0.0016 ± 0.0031) DAMA/LIBRA-1
-(0.0010 ± 0.0034) DAMA/LIBRA-2
-(0.0001 ± 0.0031) DAMA/LIBRA-3
-(0.0006 ± 0.0029) DAMA/LIBRA-4
→ statistically consistent with zero

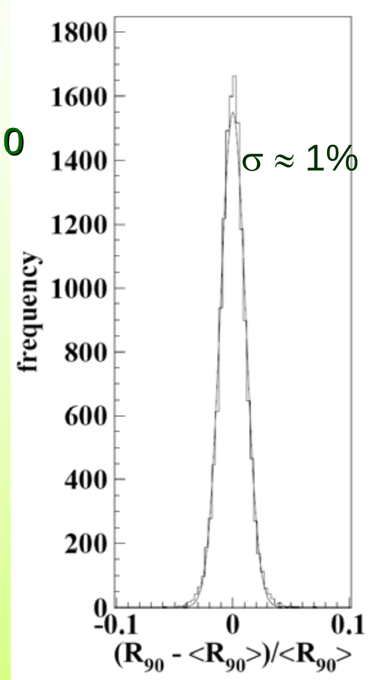
In the same energy region where the effect is observed: no modulation of the multiple-hits events (see next slide)

- **No modulation in the whole spectrum:**

studying integral rate at higher energy, R90

- R_{90} percentage variations with respect to their mean values for single crystal in the DAMA/LIBRA-1,2,3,4 running periods

→ cumulative gaussian behaviour with σ ≈ 1%, fully accounted by statistical considerations



- Fitting the behaviour with time, adding a term modulated according period and phase expected for Dark Matter particles:

Period	Mod. Ampl.
DAMA/LIBRA-1	-(0.05±0.19) cpd/kg
DAMA/LIBRA-2	-(0.12±0.19) cpd/kg
DAMA/LIBRA-3	-(0.13±0.18) cpd/kg
DAMA/LIBRA-4	(0.15±0.17) cpd/kg

consistent with zero

+ if a modulation present in the whole energy spectrum at the level found in the lowest energy region → R_{90} ~ tens cpd/kg → ~ 100 σ far away

No modulation in the background:
these results account for all sources of bckg (+ see later)

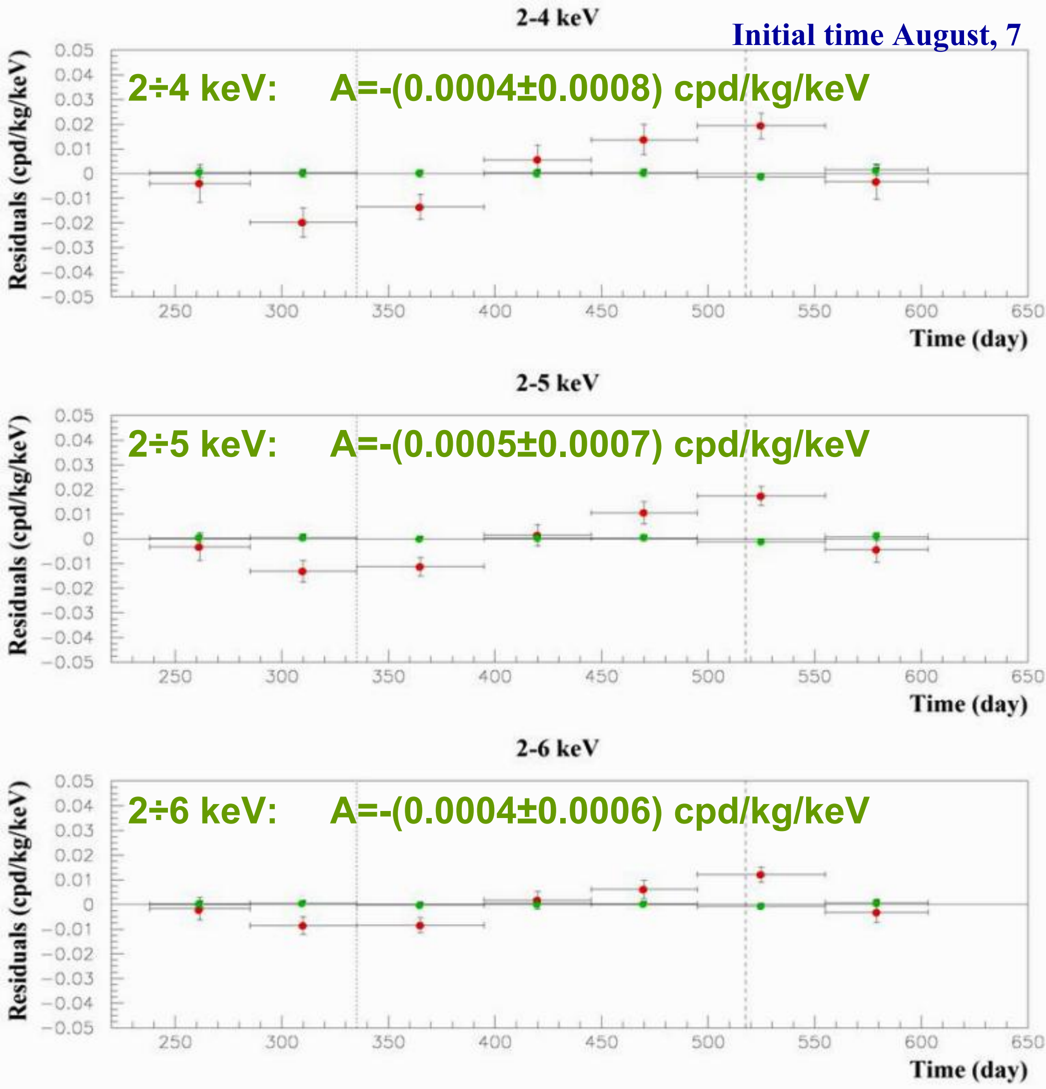
Multiple-hits events in the region of the signal - DAMA/LIBRA 1-4

- **Each detector has its own TDs read-out → pulse profiles of multiple-hits events (multiplicity > 1) acquired (exposure: 0.53 ton×yr).**
- **The same hardware and software procedures as the ones followed for single-hit events**

signals by Dark Matter particles do not belong to multiple-hits events, that is:

multiple-hits events = Dark Matter particles events “switched off”

Evidence of annual modulation with proper features as required by the DM annual modulation signature is present in the *single-hit* residuals, while it is absent in the *multiple-hits* residual rate.



2÷4 keV: A=-(0.0004±0.0008) cpd/kg/keV

2÷5 keV: A=-(0.0005±0.0007) cpd/kg/keV

2÷6 keV: A=-(0.0004±0.0006) cpd/kg/keV

This result offers an additional strong support for the presence of Dark Matter particles in the galactic halo further excluding any side effect either from hardware or from software procedures or from background

Modulation amplitudes, $S_{m,k}$, as function of the energy

The likelihood function of the *single-hit* experimental data in the *k-th* energy bin is defined as:

$$L_k = \prod_{ij} e^{-\mu_{ijk}} \frac{\mu_{ijk}^{N_{ijk}}}{N_{ijk}!}$$

N_{ijk} is the number of events collected in the *i-th* time interval (hereafter 1 day), by the *j-th* detector and in the *k-th* energy bin.

N_{ijk} follows a Poissonian distribution with expectation value:

$$\mu_{ijk} = \left[b_{jk} + R_k(t)\right] M_j \Delta t_i \Delta E \varepsilon_{jk} = \left[b_{jk} + S_{0,k} + S_{m,k} \cos\omega(t_i - t_0)\right] M_j \Delta t_i \Delta E \varepsilon_{jk}$$

The b_{jk} are the background contributions, M_j is the mass of the *j-th* detector, Δt_i is the detector running time during the *i-th* time interval, ΔE is the chosen energy bin, ε_{jk} is the overall efficiency.

The usual procedure is to minimize the function $y_k = -2\ln(L_k) - const$ for each energy bin; the free parameters of the fit are the $(b_{jk} + S_{0,k})$ contributions and the $S_{m,k}$ parameter.

The $S_{m,k}$ is the modulation amplitude of the modulated part of the signal obtained by maximum likelihood method over the data considering $T=2\pi/\omega=1$ yr and $t_0 = 152.5$ day.

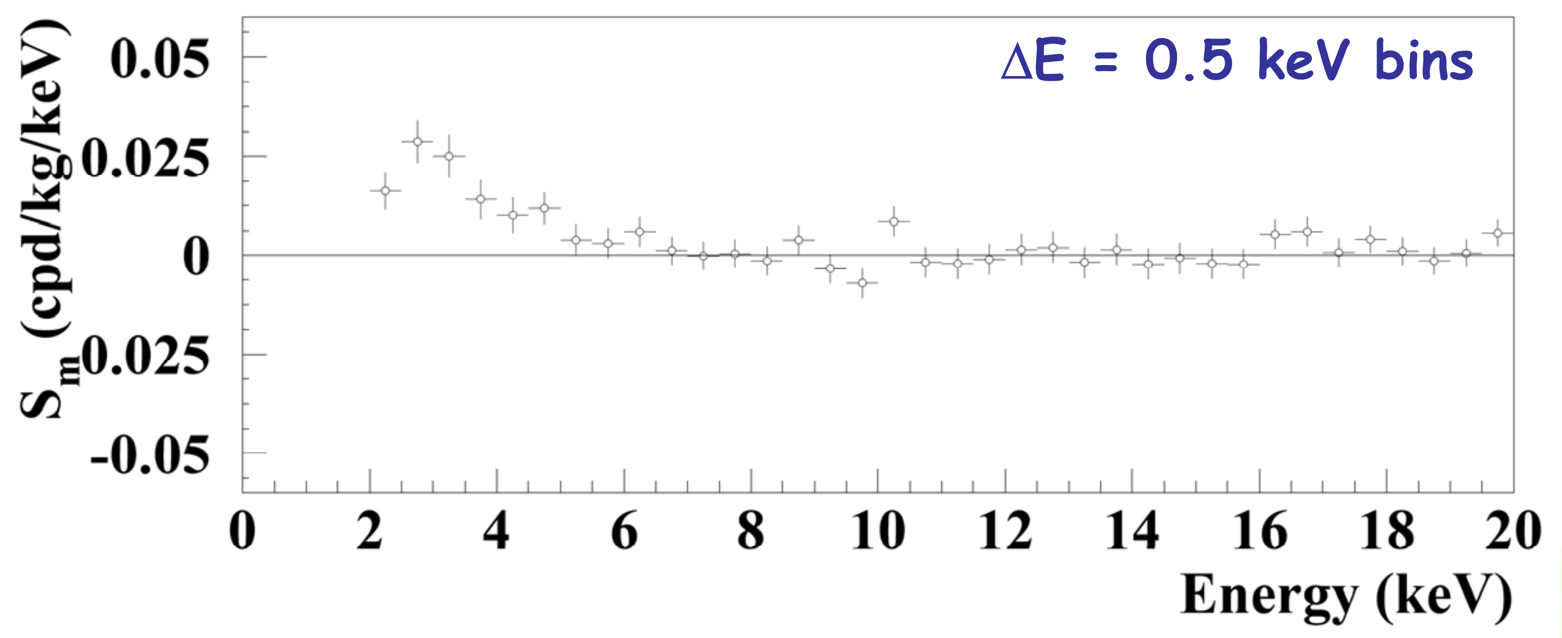
Energy distribution of the modulation amplitudes, S_m, for the total exposure

$$R(t) = S_0 + S_m \cos[\omega(t - t_0)]$$

DAMA/NaI (7 years) + DAMA/LIBRA (4 years)

total exposure: 300555 kg×day = 0.82 ton×yr

here $T=2\pi/\omega=1$ yr and $t_0= 152.5$ day



A clear modulation is present in the (2-6) keV energy interval, while S_m values compatible with zero are present just above

In fact, the S_m values in the (6–20) keV energy interval have random fluctuations around zero with χ^2 equal to 24.4 for 28 degrees of freedom

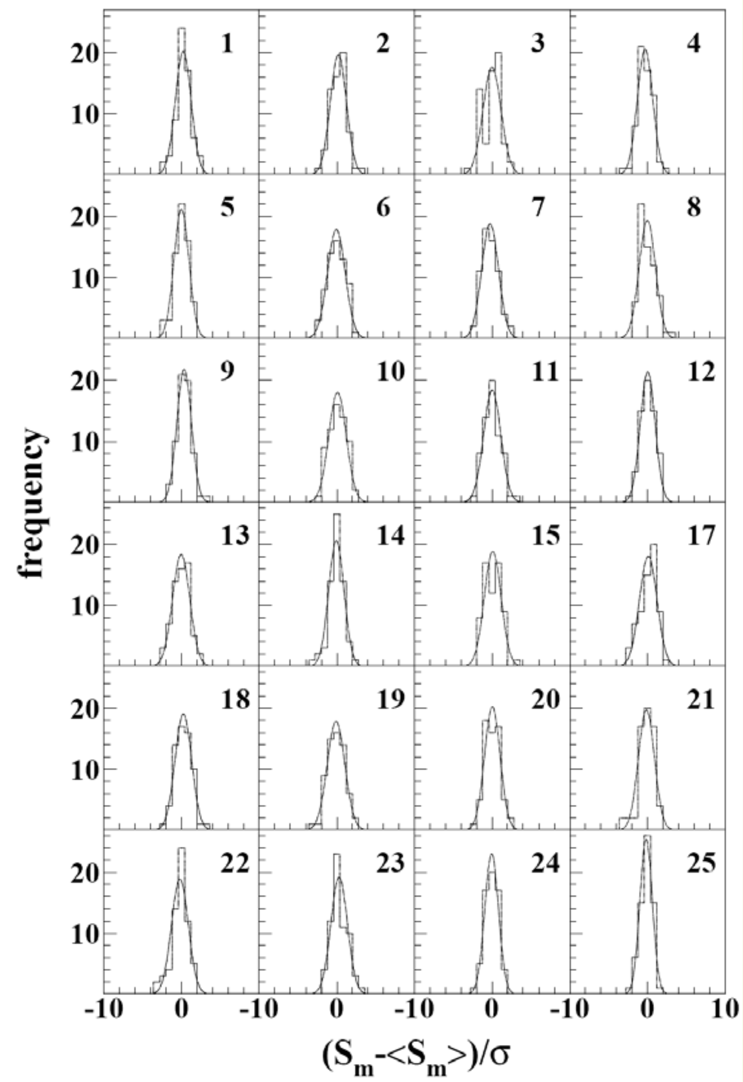
Statistical distributions of the modulation amplitudes (S_m)

a) S_m values for each detector, each annual cycle and each considered energy bin (here 0.25 keV)
b) $<S_m>$ = mean values over the detectors and the annual cycles for each energy bin; σ = errors associated to each S_m

DAMA/LIBRA (4 years)

total exposure: 0.53 ton×yr

Each panel refers to each detector separately; 64 entries = 16 energy bins in 2-6 keV energy interval × 4 DAMA/LIBRA annual cycles

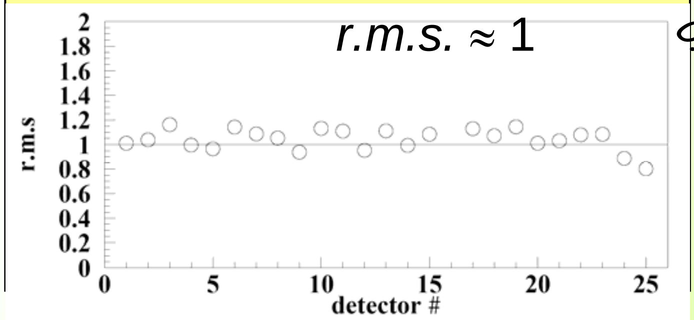


2-6 keV

Standard deviations of the variable

$(S_m-\langle S_m\rangle)/\sigma$

for the DAMA/LIBRA detectors



r.m.s. ≈ 1

0.80 < r.m.s. < 1.16

Individual S_m values follow a normal distribution since $(S_m-<S_m>)/\sigma$ is distributed as a Gaussian with a unitary standard deviation (r.m.s.)

⇒ S_m statistically well distributed in all the detectors and annual cycles

Statistical analyses about modulation amplitudes (S_m)

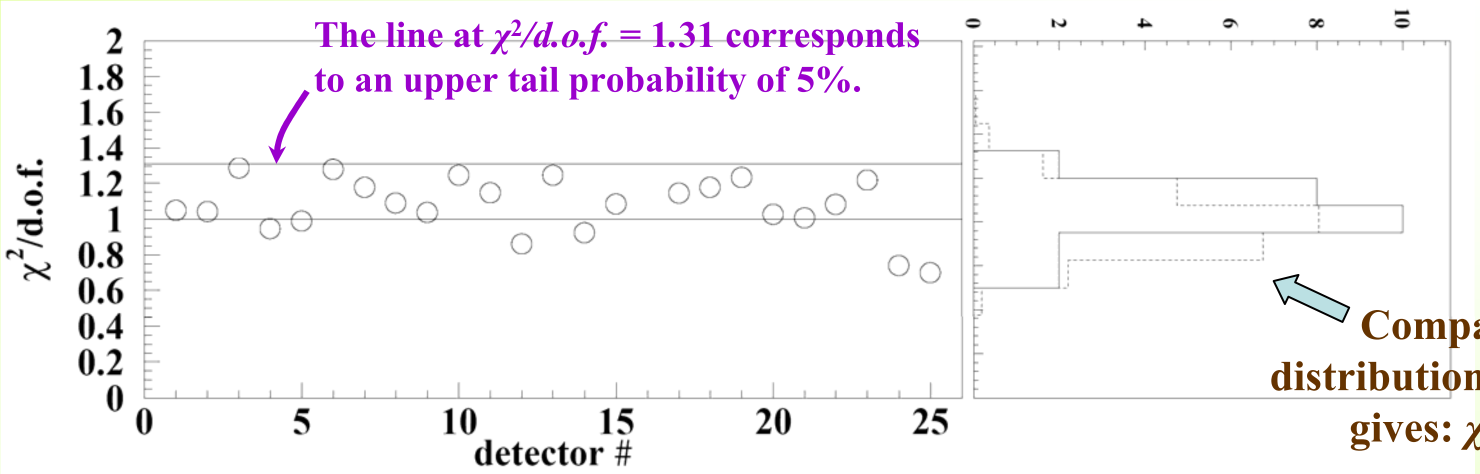
$x=(S_m-\langle S_m\rangle)/\sigma,$

$\chi^2=\Sigma x^2$

$\chi^2/d.o.f.$ values of S_m distributions for each DAMA/LIBRA detector in the (2–6) keV energy interval for the four annual cycles.

DAMA/LIBRA (4 years)

total exposure: 0.53 ton×yr



The line at $\chi^2/d.o.f. = 1.31$ corresponds to an upper tail probability of 5%.

Comparison with χ^2 distribution with 64 d.o.f. gives: $\chi^2/d.o.f. = 8.1/7$

The $\chi^2/d.o.f.$ values range from 0.7 to 1.28 (64 *d.o.f.* = 16 energy bins × 4 annual cycles)

⇒ at 95% C.L. the observed annual modulation effect is well distributed in all the detectors.

- The mean value of the twenty-four points is 1.072, slightly larger than 1. Although this can be still ascribed to statistical fluctuations, let us ascribe it to a possible systematics.
- In this case, one would have an additional error of $\leq 5\times10^{-4}$ cpd/kg/keV, if quadratically combined, or $\leq 7\times10^{-5}$ cpd/kg/keV, if linearly combined, to the modulation amplitude measured in the (2 – 6) keV energy interval.
- This possible additional error (**≤ 4.7%** or **≤ 0.7%**, respectively, of the DAMA/LIBRA modulation amplitude) can be considered as an upper limit of possible systematic effects

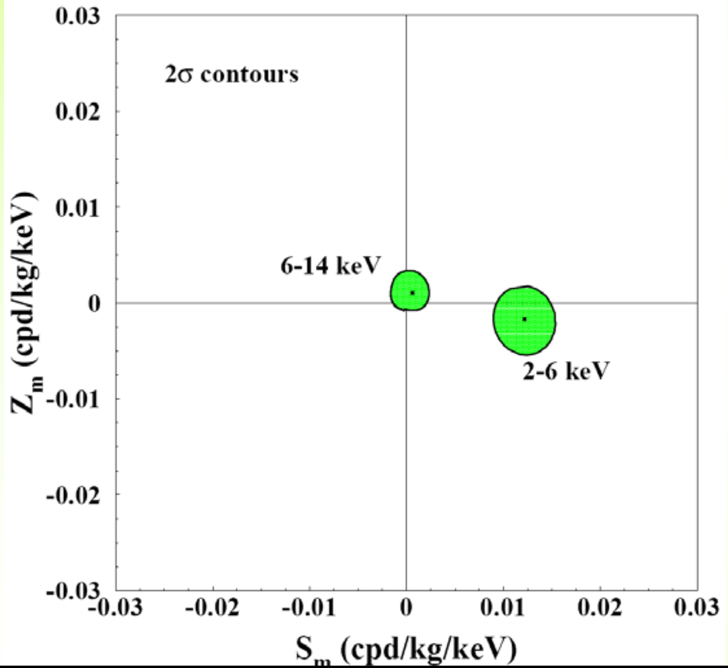
Is there a sinusoidal contribution in the signal?
Phase ≠ 152.5 day?

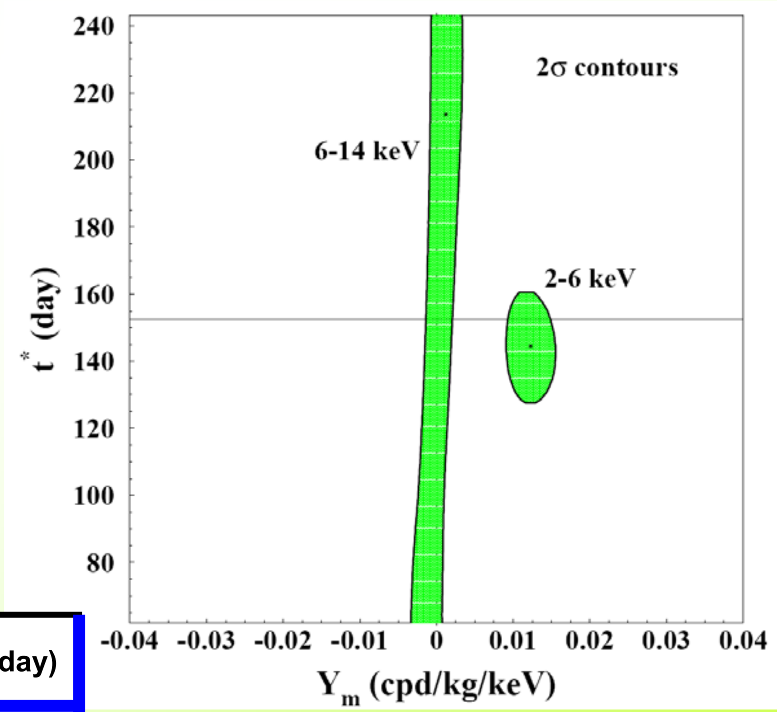
$$R(t) = S_0 + S_m \cos[\omega(t - t_0)] + Z_m \sin[\omega(t - t_0)] = S_0 + Y_m \cos[\omega(t - t^*)]$$

For Dark Matter signals:

- $|Z_m| \ll |S_m| \approx |Y_m|$
- $t^* \approx t_0 = 152.5d$
- $\omega = 2\pi/T$
- $T = 1$ year

Slight differences from 2nd June are expected in case of contributions from non thermalized DM components (as e.g. the SagDEG stream)





E (keV)	S_m (cpd/kg/keV)	Z_m (cpd/kg/keV)	Y_m (cpd/kg/keV)	t* (day)
2-6	0.0122 ± 0.0016	-0.0019 ± 0.0017	0.0123 ± 0.0016	144.0 ± 7.5
6-14	0.0005 ± 0.0010	0.0011 ± 0.0012	0.0012 ± 0.0011	--

The analysis at energies above 6 keV, the analysis of the multiple-hits events and the statistical considerations about S_m already exclude any sizeable presence of systematical effects.

Additional investigations on the stability parameters

Modulation amplitudes obtained by fitting the time behaviours of main running parameters, acquired with the production data, when including a DM-like modulation

Running conditions stable at a level better than 1%

	DAMA/LIBRA-1	DAMA/LIBRA-2	DAMA/LIBRA-3	DAMA/LIBRA-4
Temperature	-(0.0001 ± 0.0061) °C	(0.0026 ± 0.0086) °C	(0.001 ± 0.015) °C	(0.0004 ± 0.0047) °C
Flux N_2	(0.13 ± 0.22) l/h	(0.10 ± 0.25) l/h	-(0.07 ± 0.18) l/h	-(0.05 ± 0.24) l/h
Pressure	(0.015 ± 0.030) mbar	-(0.013 ± 0.025) mbar	(0.022 ± 0.027) mbar	(0.0018 ± 0.0074) mbar
Radon	-(0.029 ± 0.029) Bq/m^3	-(0.030 ± 0.027) Bq/m^3	(0.015 ± 0.029) Bq/m^3	-(0.052 ± 0.039) Bq/m^3
Hardware rate above single photoelectron	-(0.20 ± 0.18) × 10^{-2} Hz	(0.09 ± 0.17) × 10^{-2} Hz	-(0.03 ± 0.20) × 10^{-2} Hz	(0.15 ± 0.15) × 10^{-2} Hz

All the measured amplitudes well compatible with zero

+none can account for the observed effect

(to mimic such signature, spurious effects and side reactions must not only be able to account for the whole observed modulation amplitude, but also simultaneously satisfy all the 6 requirements)

Summary of the results obtained in the additional investigations of possible systematics or side reactions

(DAMA/LIBRA – NIMA592(2008)297, EPJC56(2008)333)

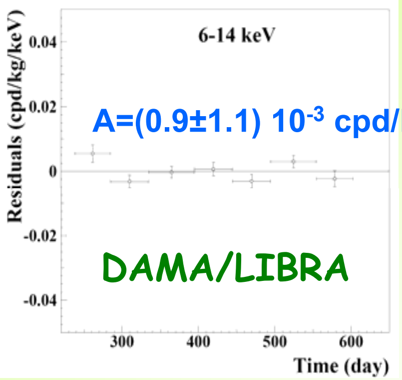
Source	*Main comment*	*Cautious upper limit (90%C.L.)*
RADON	Sealed Cu box in HP Nitrogen atmosphere, 3-level of sealing, etc.	**$<2.5\times10^{-6}$ cpd/kg/keV**
TEMPERATURE	Installation is air conditioned+ detectors in Cu housings directly in contact with multi-ton shield→ huge heat capacity + T continuously recorded	**$<10^{-4}$ cpd/kg/keV**
NOISE	Effective full noise rejection near threshold	**$<10^{-4}$ cpd/kg/keV**
ENERGY SCALE	Routine + instrinsic calibrations	**$<1\text{-}2\times10^{-4}$ cpd/kg/keV**
EFFICIENCIES	Regularly measured by dedicated calibrations	**$<10^{-4}$ cpd/kg/keV**
BACKGROUND	No modulation above 6 keV; no modulation in the (2-6) keV *multiple-hits* events; this limit includes all possible sources of background	**$<10^{-4}$ cpd/kg/keV**
SIDE REACTIONS	Muon flux variation measured by MACRO	**$<3\times10^{-5}$ cpd/kg/keV**

+ even if larger they cannot satisfy all the requirements of annual modulation signature

Thus, they can not mimic the observed annual modulation effect

Summarizing on a hypothetical background modulation in DAMA/LIBRA 1-4

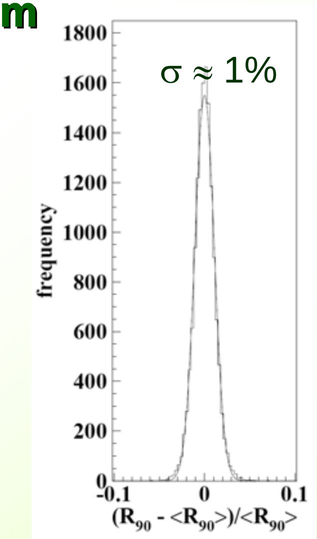
- **No Modulation above 6 keV**



$A=(0.9\pm1.1)\ 10^{-3}$ cpd/kg/keV

DAMA/LIBRA

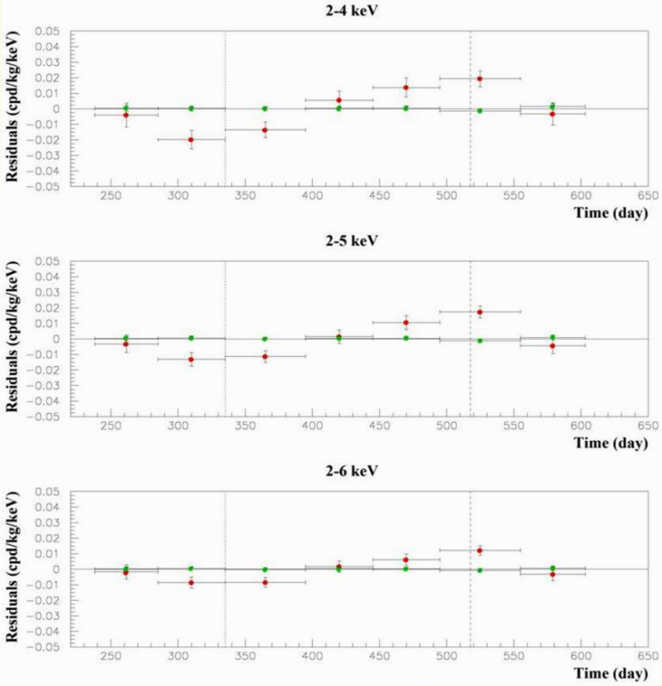
- **No modulation in the whole energy spectrum**



$\sigma \approx 1\%$

+ if a modulation present in the whole energy spectrum at the level found in the lowest energy region → R_{90} ~ tens cpd/kg → ~ 100 σ far away

- **No modulation in the 2-6 keV *multiple-hits* residual rate**



multiple-hits residual rate (green points) vs single-hit residual rate (red points)

No background modulation (and cannot mimic the signature):
all this accounts for the all possible sources of bckg

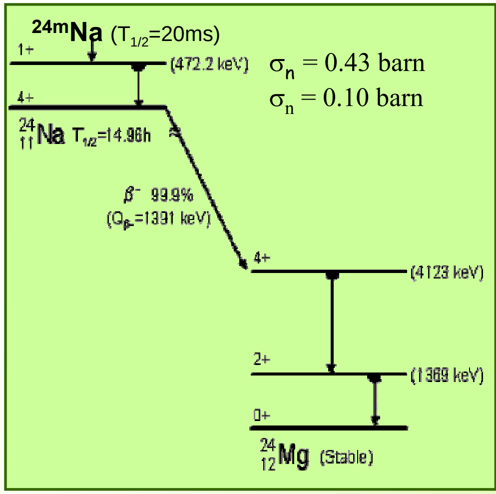
Nevertheless, additional investigations performed ...

Can a possible thermal neutron modulation account for the observed effect?

NO

- Thermal neutrons flux measured at LNGS :

$\Phi_n = 1.08\ 10^{-6}$ n cm^{-2} s^{-1} (N.Cim.A101(1989)959)



σ_n = 0.43 barn

σ_n = 0.10 barn

- Experimental upper limit on the thermal neutrons flux "*surviving*" the neutron shield in DAMA/LIBRA:
  - studying triple coincidences able to give evidence for the possible presence of ^{24}Na from neutron activation:

$\Phi_n < 1.2 \times 10^{-7}$ n cm^{-2} s^{-1} (90%C.L.)

- Two consistent upper limits on thermal neutron flux have been obtained with DAMA/NaI considering the same capture reactions and using different approaches.

Evaluation of the expected effect:

- Capture rate = $\Phi_n\ \sigma_n\ N_T < 0.022$ captures/day/kg

HYPOTHESIS: assuming very cautiously a 10% thermal neutron modulation:

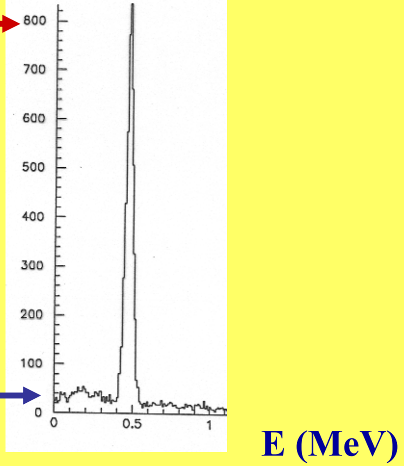
$S_m^{(\text{thermal n})} < 0.8 \times 10^{-6}$ cpd/kg/keV ($< 0.01\%\ S_m^{\text{observed}}$)

MC simulation of the process

When $\Phi_n = 10^{-6}$ n cm^{-2} s^{-1}:

$7 \cdot 10^{-5}$ cpd/kg/keV

$1.4 \cdot 10^{-3}$ cpd/kg/keV



In all the cases of neutron captures (^{24}Na, ^{128}I, ...) a possible thermal n modulation induces a variation in all the energy spectrum

Already excluded also by R_{90} analysis

Can the μ modulation measured by MACRO account for the observed effect?

Case of fast neutrons produced by muons

Φ_μ @ LNGS ≈ 20 $\mu\ m^{-2}\ d^{-1}$ **(±2% modulated)**

Neutron Yield @ LNGS: $Y = 1 \div 7\ 10^{-4}\ n/\mu/(g/cm^2)$ **(hep-ex/0006014)**

R_n = (fast n by μ)/(time unit) = $\Phi_\mu\ Y\ M_{eff}$

Annual modulation amplitude at low energy due to μ modulation:

$$S_m^{(\mu)} = R_n\ g\ \varepsilon\ f_{\Delta E}\ f_{single}\ 2\% / (M_{setup}\ \Delta E)$$

where:

g = geometrical factor

ε = detection efficiency by elastic scattering

$f_{\Delta E}$ = energy window (E>2keV) efficiency

f_{single} = single hit efficiency

Hyp.: M_{eff} = 15 tons

$g \approx \varepsilon \approx f_{\Delta E} \approx f_{single} \approx 0.5$ (cautiously)

Knowing that:

$M_{setup} \approx 250$ kg and ΔE=4keV

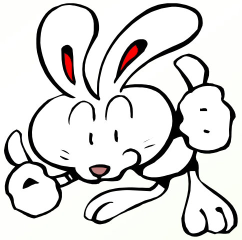
$S_m^{(\mu)} < (0.4 \div 3) \times 10^{-5}$ cpd/kg/keV

NO

Moreover, this modulation also induces a variation in other parts of the energy spectrum + different phase

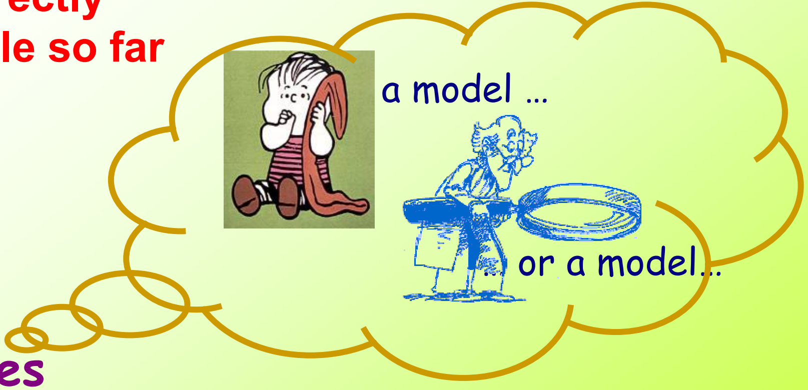
It cannot mimic the signature: already excluded also by R_{90}

The positive and model independent result by DAMA/NaI + DAMA/LIBRA

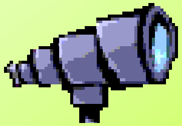


- Presence of modulation for 11 annual cycles at ~8.2σ C.L. with the proper distinctive features of the DM signature; all the features satisfied by the data over 11 independent experiments of 1 year each one
- Absence of known sources of possible systematics and side processes able to quantitatively account for the observed modulation amplitude and to contemporaneously satisfy all the peculiarities of the signature

No other experiment whose result can be directly compared in model independent way is available so far



Corollary quests for candidates



Model-independent evidence by DAMA/NaI and DAMA/LIBRA

well compatible with several candidates (in several of the many astrophysical, nuclear and particle physics scenarios); other ones are open

Neutralino as LSP in SUSY theories

Various kinds of WIMP candidates with several different kind of interactions Pure SI, pure SD, mixed + Migdal effect +channeling,… (from low to high mass)

a heavy ν of the 4-th family

Pseudoscalar, scalar or mixed light bosons with axion-like interactions

WIMP with preferred inelastic scattering

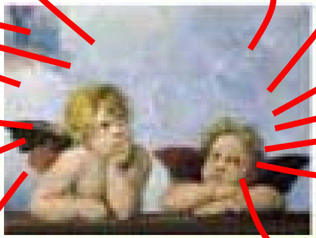
Mirror Dark Matter

Light Dark Matter

Self interacting Dark Matter

Dark Matter (including some scenarios for WIMP) electron-interacting

Sterile neutrino



heavy exotic canditates, as “4th family atoms”, ...

Elementary Black holes such as the Daemons

Kaluza Klein particles

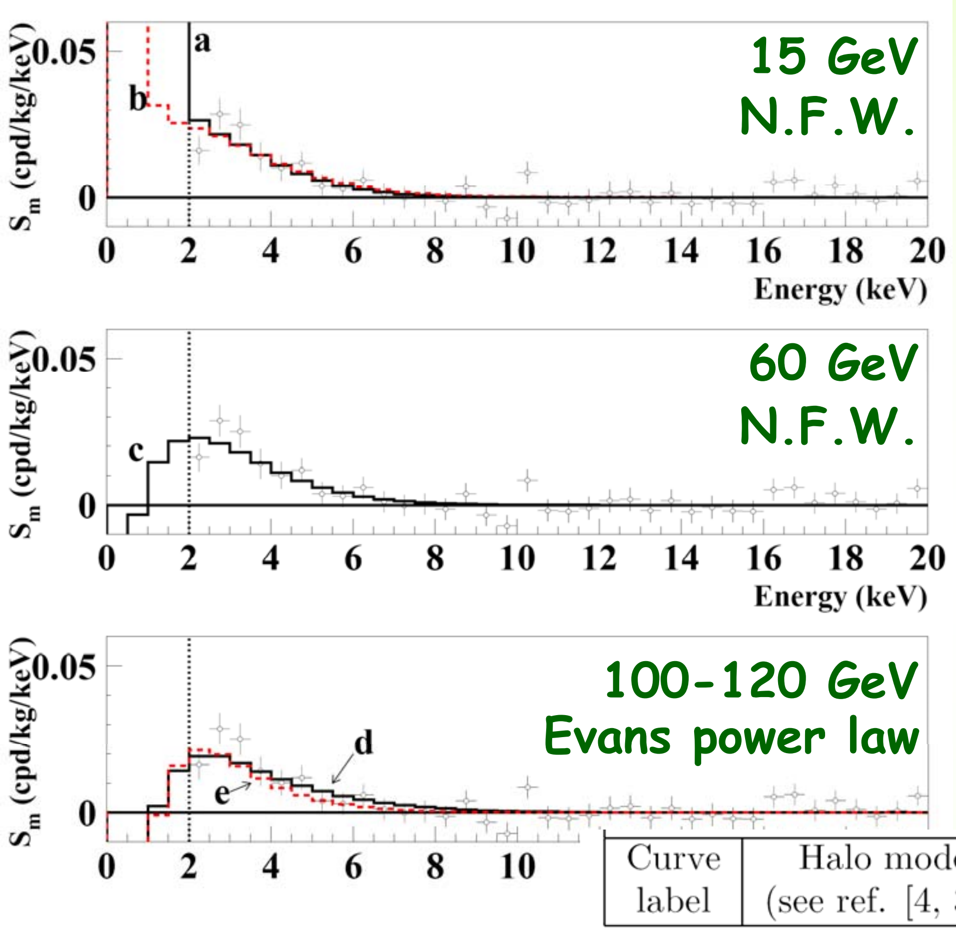
… and more

Possible model dependent positive hints from indirect searches not in conflict with DAMA results

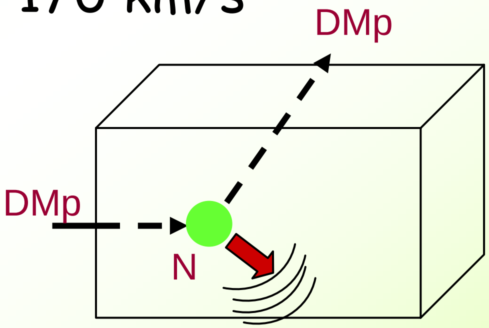
(but interpretation, evidence itself, derived mass and cross sections depend e.g. on bckg modeling, on DM spatial velocity distribution in the galactic halo, etc.)

Available results from direct searches using different target materials and approaches do not give any robust conflict

Examples for few of the many possible scenarios superimposed to the measured modulation amplitues $S_{m,k}$



WIMP DM candidate (as in [4])
considering elastic scattering on nuclei
SI dominant coupling
v_0 = 170 km/s



About the same C.L.
(no best fit values)

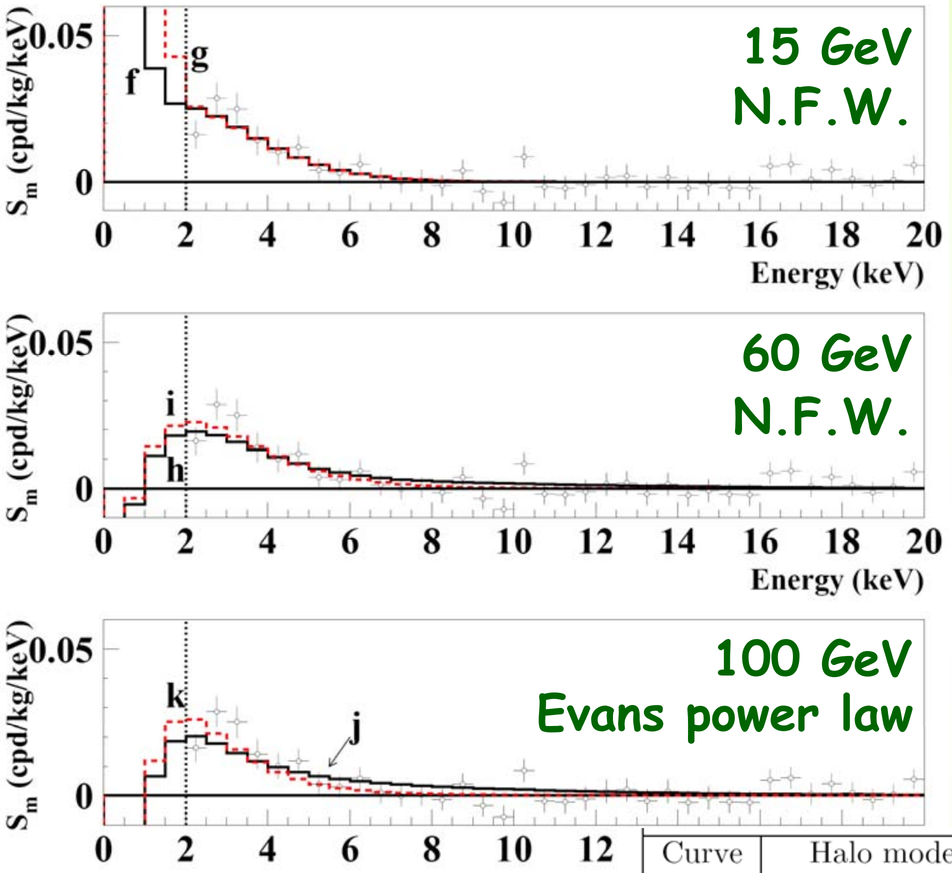
…scaling from NaI

Curve label	Halo model (see ref. [4, 34])	Local density (GeV/cm^3)	Set as in [4]	DM particle mass	$\xi\sigma_{SI}$ (pb)
a	A5 (NFW)	0.2	A	15 GeV	3.1×10^{-4}
b	A5 (NFW)	0.2	A	15 GeV	1.3×10^{-5}
c	A5 (NFW)	0.2	B	60 GeV	5.5×10^{-6}
d	B3 (Evans power law)	0.17	B	100 GeV	6.5×10^{-6}
e	B3 (Evans power law)	0.17	A	120 GeV	1.3×10^{-5}

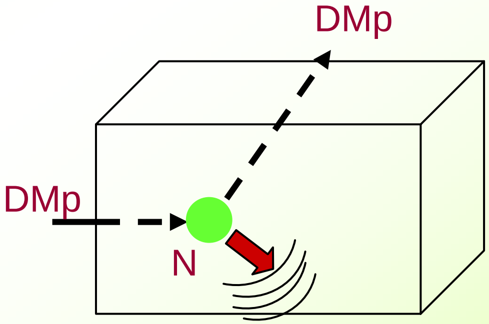
channeling contribution as in EPJC53(2008)205 considered for curve *b*

[4] RNC 26 (2003) 1; [34] PRD66 (2002) 043503

Examples for few of the many possible scenarios superimposed to the measured modulation amplitues $S_{m,k}$



WIMP DM candidate (as in [4])
Elastic scattering on nuclei
SI & SD mixed coupling
v_0 = 170 km/s



About the same C.L.
(no best fit values)

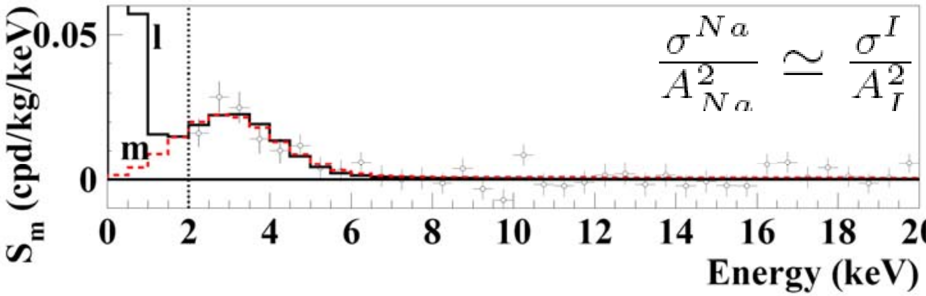
…scaling from NaI

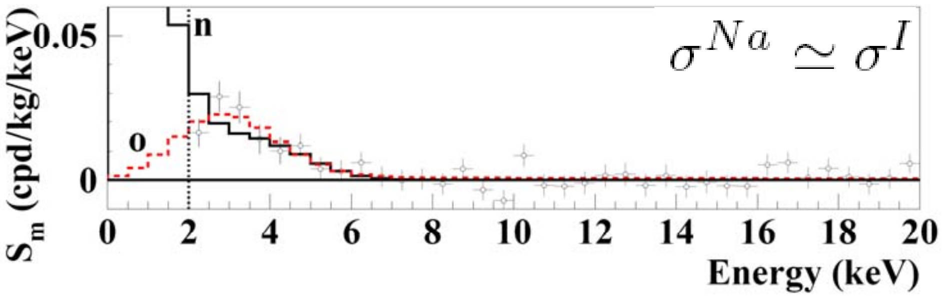
$$\theta = 2.435$$

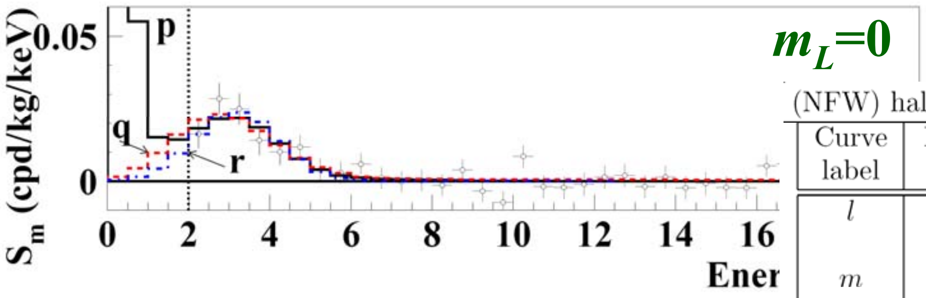
Curve label	Halo model (see ref. [4, 34])	Local density (GeV/cm^3)	Set as in [4]	DM particle mass	$\xi\sigma_{SI}$ (pb)	$\xi\sigma_{SD}$ (pb)
f	A5 (NFW)	0.2	A	15 GeV	10^{-7}	2.6
g	A5 (NFW)	0.2	A	15 GeV	1.4×10^{-4}	1.4
h	A5 (NFW)	0.2	B	60 GeV	10^{-7}	1.4
i	A5 (NFW)	0.2	B	60 GeV	8.7×10^{-6}	8.7×10^{-2}
j	B3 (Evans power law)	0.17	A	100 GeV	10^{-7}	1.7
k	B3 (Evans power law)	0.17	A	100 GeV	1.1×10^{-5}	0.11

[4] RNC 26 (2003) 1; [34] PRD66 (2002) 043503

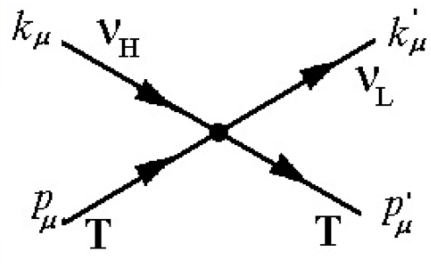
Examples for few of the many possible scenarios superimposed to the measured modulation amplitues $S_{m,k}$



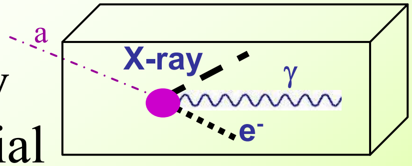




LDM candidate (as in arXiv:0802.4336): inelastic interaction with electron or nucleus targets



Light bosonic candidate (as in IJMPA21(2006)1445): axion-like particles totally absorbed by target material



About the same C.L. (no best fit values)

(NFW) halo model as in [4, 34], local density = 0.17 GeV/cm³, local velocity = 170 km/s

Curve label	DM particle	Interaction	Set as in [4]	m_H	Δ	Cross section (pb)
l	LDM	coherent on nuclei	A	30 MeV	18 MeV	$\xi\sigma_m^{coh} = 1.8 \times 10^{-6}$
m	LDM	coherent on nuclei	A	100 MeV	55 MeV	$\xi\sigma_m^{coh} = 2.8 \times 10^{-6}$
n	LDM	incoherent on nuclei	A	30 MeV	3 MeV	$\xi\sigma_m^{inc} = 2.2 \times 10^{-2}$
o	LDM	incoherent on nuclei	A	100 MeV	55 MeV	$\xi\sigma_m^{inc} = 4.6 \times 10^{-2}$
p	LDM	coherent on nuclei	A	28 MeV	28 MeV	$\xi\sigma_m^{coh} = 1.6 \times 10^{-6}$
q	LDM	incoherent on nuclei	A	88 MeV	88 MeV	$\xi\sigma_m^{inc} = 4.1 \times 10^{-2}$
r	LDM	on electrons	–	60 keV	60 keV	$\xi\sigma_m^{e} = 0.3 \times 10^{-6}$

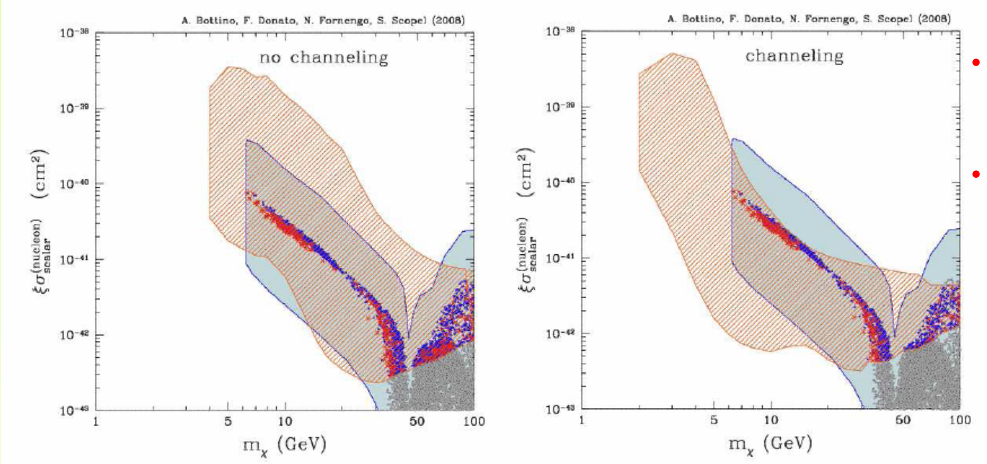
curve *r*: also pseudoscalar axion-like candidates (e.g. majoron)
m_a=3.2 keV g_{aee}= 3.9 10^{-11}

[4] RNC 26 (2003) 1; [34] PRD66 (2002) 043503

... some examples appeared in literature...

Supersymmetric expectations in MSSM

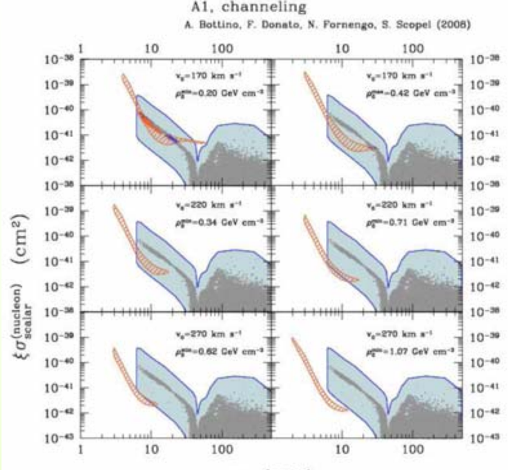
A. Bottino et al., PRD78 (2008) 083520



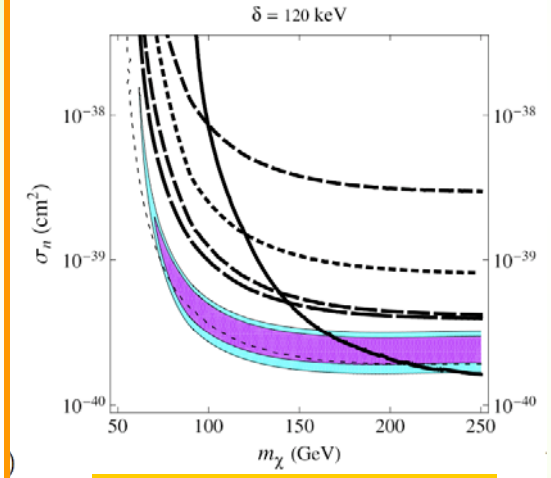
- *Assuming for the neutralino a dominant purely SI coupling*
- *when releasing the gaugino mass unification at GUT scale: $M_1/M_2 \neq 0.5$ (<);*

(where M_1 and M_2 U(1) and SU(2) gaugino masses)

low mass configurations are obtained



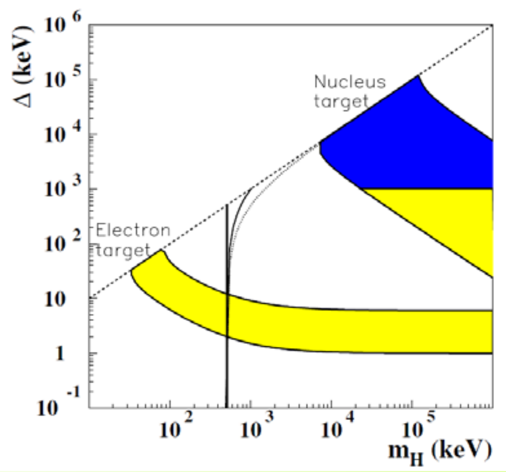
Inelastic DM



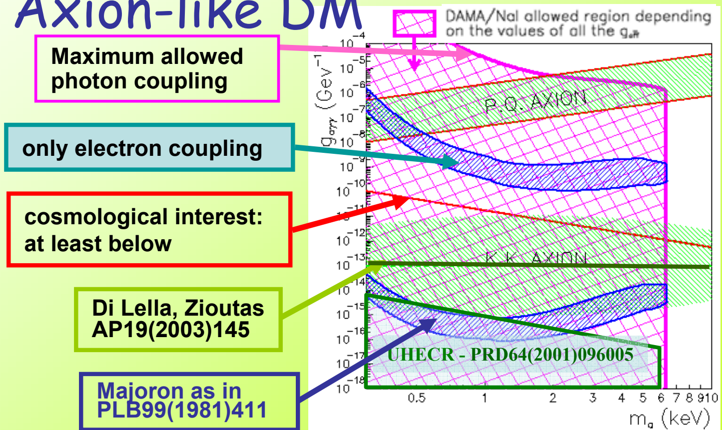
Weiner&Tucker-Smith

Light DM

MPLA23 (2008) 2125



Axion-like DM



Maximum allowed photon coupling

only electron coupling

cosmological interest: at least below

Di Lella, Zioutas AP19(2003)145

Majoron as in PLB99(1981)411

where we are …

- DAMA/LIBRA over 4 annual cycles (0.53 ton×yr) confirms the results of DAMA/NaI (0.29 ton×yr)
- The cumulative confidence level for the model independent evidence for presence of DM particle in the galactic halo is 8.2 σ (total exposure 0.82 ton × yr)

- **First upgrading of the experimental set-up in Sept. 2008**

Phase 1

- Mounting of the “clean room” set-up in order to operate in HP N_2 atmosphere
- Opening of the shield of DAMA/LIBRA set-up in HP N_2 atmosphere
- Replacement of some PMTs in HP N_2 atmosphere
- Closing of the shield

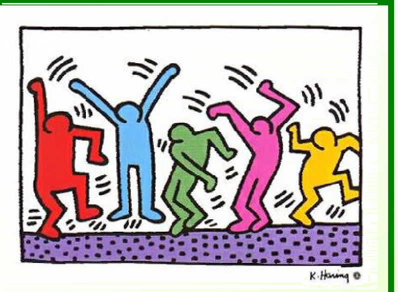
Phase 2

- Dismounting of the Tektronix TDs (Digitizers + Crates)
- Mounting of the new Acqiris TD (Digitizers + Crate)
- Mounting of the new DAQ system with optical read-out
- Test of the new TDs (*hardware*) and of the new required DAQ system *(software*)

… since Oct. 2008 again in data taking

… and where DAMA is going to

- Continuing the data taking
- Update corollary analyses in some of the many possible scenarios for DM candidates, interactions, halo models, nuclear/atomic properties, etc..



- ***Next upgrading***: replacement of all the PMTs with higher Q.E. ones.
- Production of new high Q.E. PMTs in progress
- Goals:
  - better separation under 2 keV in the rejection plane between noise and single-hit scintillation events
  - lowering the energy threshold (presently, at 2 keV)
  - improvement of the acceptance efficiency near energy threshold
  - increase of the sensitivity in the *model independent* analysis (amplitude, phase, second order effects, …)
  - improvement of the sensitivity in the *model dependent* analyses, allowing to better disentangle several astrophysical, particle physics and nuclear physics scenarios

- Analyses/data taking to investigate also other rare processes in progress/foreseen

- Long term data taking to improve the investigation, to disentangle at least some of the many possibilities, to investigate other features of DM particle component(s), second order effects, etc..

A possible highly radiopure NaI(Tl) multi-purpose set-up DAMA/1 ton (proposed by DAMA in 1996) at R&D phase

to deep investigate Dark Matter phenomenology at galactic scale

Felix qui potuit rerum cognoscere causas *(Virgilio, Georgiche, II, 489)*